

## Recent Results from Studies of Electric Discharges in the Mesosphere

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**Abstract** The paper reviews recent advances in studies of electric discharges in the stratosphere and mesosphere above thunderstorms, and their effects on the atmosphere. The primary focus is on the sprite discharge occurring in the mesosphere, which is the most commonly observed high altitude discharge by imaging cameras from the ground, but effects on the upper atmosphere by electromagnetic radiation from lightning are also considered. During the past few years, co-ordinated observations over Southern Europe have been made of a wide range of parameters related to sprites and their causative thunderstorms. Observations have been complemented by the modelling of processes ranging from the electric discharge to perturbations of trace gas concentrations in the upper atmosphere. Observations point to significant energy deposition by sprites in the neutral atmosphere as observed by infrasound waves detected at up to 1000 km distance, whereas elves and lightning have been shown significantly to affect ionization and heating of the lower ionosphere/mesosphere. Studies of the thunderstorm systems powering high altitude discharges show the important role of intracloud (IC) lightning in sprite generation as seen by the first simultaneous observations of IC activity, sprite activity and broadband, electromagnetic radiation in the VLF range. Simulations of sprite ignition suggest that, under certain conditions, energetic electrons in the runaway regime are generated in streamer discharges. Such electrons may be the source of X- and Gamma-rays observed in lightning, thunderstorms and the so-called Terrestrial Gamma-ray Flashes (TGFs) observed from space over thunderstorm regions. Model estimates of sprite perturbations to the global atmospheric electric circuit, trace gas concentrations and atmospheric dynamics suggest significant local perturbations, and possibly significant meso-scale effects, but negligible global effects.

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## 1 Introduction

It is less than 20 years since scientists first documented spectacular optical flashes in the stratosphere and mesosphere above large thunderstorm systems (Franz et al. 1990). The flashes were discovered over the US plains, where summer thunderstorms are powered by hot humid air masses sweeping up from the Caribbean and the Pacific. The discovery was a confirmation of predictions made by the Nobel prize winner C.T.R. Wilson in the 1920s (Wilson 1925), and led researchers to examine the literature of the past, discovering anecdotal evidence dating as far back as the 19th century. A careful re-examination of such data sources as video recordings made by the space shuttle payload bay cameras further identified previously unnoticed emissions over Africa and Australia (Vaughan and Vonnegut 1989; Boeck et al. 1995; Boeck et al. 1998). In the years following, ground and aircraft campaigns were conducted, mostly in the US, but also in Australia, Japan, Taiwan, South America, and Europe. In all regions the occurrence of optical emissions in the stratosphere and mesosphere associated with thunderstorms was documented.

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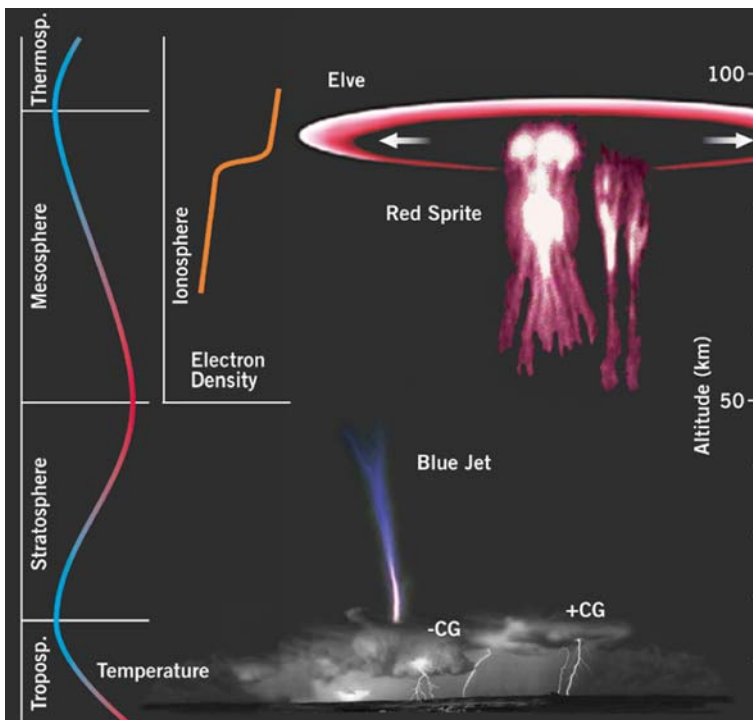
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Research in the past decade has identified a surprising variety of emissions (Neubert 2003). The most commonly observed emission from ground, the so-called “sprite”, is a manifestation of electrical breakdown of the mesosphere at 40–80 km altitude (Sentman et al. 1995). Sprites are almost always associated with positive cloud-to-ground (+CG) lightning discharges which lower positive charge from clouds to the ground (Boccippio et al. 1995). A CG lightning discharge, or flash, is composed of several strokes, each involving a downward leader and an upward return stroke (Rakov and Uman 2003). The “blue jet” is a discharge propagating upwards into the stratosphere from cloud tops. It is the discharge that most resembles classical lightning except that it propagates upwards into the stratosphere, to about 40 km altitude. It may or may not be associated with cloud-to-ground (CG) lightning activity (Wescott et al. 1995). Then there is the “elve”, which is a concentric ring of optical emissions propagating horizontally outwards at the bottom edge of ionosphere at  $\sim 90$  km altitude (Fukunishi et al. 1996). Elves are associated with CG lightning of either polarity and are caused by the electromagnetic pulse radiated by the CG discharge current (Inan et al. 1996; Cho and Rycroft 1998; Füllekrug et al. 2006a). Observations have further documented the “gigantic jet”, a discharge where a blue jet triggers a sprite, creating electrical breakdown of the atmosphere from the thunderstorm clouds directly up to the bottom of the ionosphere (Pasko et al. 2002; Su et al. 2003). The panoply of the most commonly observed emissions of the upper atmosphere, also referred to as “Transient Luminous Events” (TLEs), is shown in Fig. 1.

Within the past decade, a different, but possibly related, type of flash has been discovered. It is the Terrestrial Gamma-ray Flash (TGF) from the atmosphere above



**Fig. 1** Diagram illustrating the variety of upper atmospheric flashes and the altitudes at which they occur (Neubert 2003)

thunderstorms, first recorded by the Compton Gamma Ray Observatory (CGRO) satellite (Fishman et al. 1994; Nemiroff et al. 1997; Østgaard et al. 2008). More recent observations come from the Reuven Ramaty High Energy Solar Spectroscopic Imager (RHessi), detecting 10–20 TGFs per month with energies up to 20 MeV (Smith et al. 2005). It has been suggested that TGFs are bremsstrahlung radiation from upward propagating, relativistic (MeV) electron beams generated in a runaway discharge process powered by the transient electric field in the stratosphere and mesosphere following a lightning event. The runaway discharge process has further been suggested for the initiation of lightning and sprites (Roussel-Dupré and Gurevich 1996), but there is no evidence yet for a direct connection between sprites and TGFs.

In this paper we discuss the most recent results of European research into electric discharges in the mesosphere. The research was conducted within the Research Training Network (RTN) “Coupling of Atmospheric Layers” (2002–2006) of 11 participating research institutions and 9 young doctoral students and post-doctoral fellows. The network fielded annual observational campaigns to Southern Europe, the *EuroSprite* 2003–2006 campaigns, and developed computational tools to model sprite discharges and their various effects on the mesosphere. Emphasis in this paper is placed on sprites, which are the most commonly observed high-altitude discharges in the campaigns.

### 1.1 The Electric Earth

Up to  $\sim 2000$  thunderstorms are active over the surface of the Earth at any given time, with an average of 40–50 lightning discharges occurring every second (Christian et al. 2003). The most active regions of lightning are over the Americas, Africa, and South-East Asia, with the major proportion of lightning being over land. Lightning discharges radiate intense electromagnetic (EM) pulses, from a few Hz to hundreds of MHz, of  $\sim 20$  GW peak power for  $\sim 1$  ms to  $\sim 1$  s duration, as measured by electric and magnetic field sensors. The EM radiation heats the partly ionized layers of the upper atmosphere called the D- and E-regions (60–120 km) of the ionosphere. Lightning also produces quasi-static electric fields up to 1 kV/m at mesospheric altitudes, which are capable of accelerating electrons to relativistic energies.

In general, thunderclouds charge to positive potentials at their tops and negative potentials at their bottoms. The global distribution of thunderstorms generates an electric current from the top of the clouds to the ionosphere, charging the ionosphere to  $\sim 250$  kV relative to the Earth’s surface. In fair weather regions, the electric field of the atmosphere at ground level is 120–130 V/m pointing downwards, and the electric current density in the air from the ionosphere to ground is  $\sim 2$  pAm<sup>-2</sup> (Williams 2002). This simplistic picture does not describe the heterogeneous character of the charge distributions found in thunderclouds, with many charge-carrying regions of varying polarity. While in situ measurements of electric fields within thunderclouds and above have been made from rockets and balloons (Stolzenburg et al. 2007), they are rare because of the difficulty of such measurements. The fields at mesospheric altitudes where sprites occur are primarily estimated from simple assumptions on the charge distributions in the thunderclouds and from observations of the AC components of the electromagnetic field radiated by lightning discharges.

Sparked by the discovery of TLEs and TGFs, our understanding of fundamental processes related to atmospheric electricity is under re-evaluation (Arnold and Neubert 2002). It is now appreciated that thunderstorms of the troposphere couple in important ways to the stratosphere, mesosphere and near-Earth space. It has been proposed that electric currents in the atmosphere affect cloud formation through the electro-scavenging process, where

electric charges on ice particles and electric double layers in clouds enhance the growth of ice crystals (Tinsley et al. 2000, 2001). Lightning also generates Very Low Frequency (VLF) electromagnetic radiation that propagates through the magnetosphere in the whistler mode. These waves scatter energetic electrons from the radiation belts to precipitate into the atmosphere (Rycroft 1973) where they can ionize the atmosphere above  $\sim 60$  km altitude (Inan et al. 2007). Finally, the strong convection of thunderstorms powers the transport of water into the upper troposphere and through the tropopause to the stratosphere, where this important greenhouse gas is distributed over wider regions (Chaboureau et al. 2007), as well as gravity waves that propagate to the mesosphere where they break and deposit their energy (Dewan et al. 1998).

Processes in space also couple to the atmosphere, perturbing the electric environment. Solar X-ray and UV radiation ionize the upper atmosphere creating the ionosphere with high electrical conductivity, and energetic charged particles from the Sun and the magnetosphere generate bursts of enhanced ionization in the mesosphere and ionosphere. Electric fields and currents generated by solar wind-magnetosphere interactions affect the ionosphere and thermosphere. Cosmic rays control the electrical conductivity in most regions below the mesosphere, and may create the seed charges for the runaway electron discharge mechanism mentioned earlier. It is also suggested that galactic cosmic rays influence the production of ion-mediated cloud condensation nuclei, and thus cloud formation (Svensmark et al. 2007). Research efforts in atmospheric electrical processes further support the view, also put forward in other areas of atmospheric research, that global models of atmospheric circulation, properly representing the dynamics, must include all atmospheric layers from the surface to the thermosphere and ionosphere.

While TLEs are understood in general terms, many aspects of their behaviour are not known, including their chemical effects, which are virtually unexplored. Modification of trace gas concentrations such as the oxides of nitrogen ( $\text{NO}_x$ ) and ozone could be important as they affect the Earth's radiative balance. While the lifetime of these gases is relatively short in the troposphere (weeks), the gases may reside for years in the mesosphere. In addition, while processes in the middle and upper atmosphere, such as stratospheric cooling associated with global warming and solar-terrestrial events, are well known, the coupling to the troposphere and the overall dynamic response of the atmosphere are not understood.

The mesosphere and lower thermosphere are the regions of the atmosphere about which the least is known. They are difficult to access as they are too low for in situ spacecraft observations, too high for balloon observations, and remote sensing is hampered by the low densities of important trace species and the high degree of variability over a broad range of temporal and spatial scales. As a consequence, these regions have received less attention than more readily accessible regions, both above and below. In many respects, the mesosphere and lower thermosphere represent an uncharted “frontier” for upper atmospheric science. TLEs and TGFs represent a new coupling path between the troposphere, stratosphere, mesosphere, and the near-Earth space environment. An in-depth understanding of TLEs and TGFs holds the additional promise of improving our knowledge of the regions in which they occur.

## 1.2 Red Sprites: Observations

Red sprites are large but brief luminous flashes that appear directly above an active thunderstorm system—usually coincident with a positive cloud-to-ground stroke (Boccippio et al. 1995). They may occur in clusters reaching horizontal distances of

50–100 km, or as single luminous columns (c-sprites) (Wescott et al. 1998). The brightest region is in the altitude range 65–85 km with most of the intensity in the red, and with blue tendril-like filamentary structures extending downward at times to as low as the cloud top. The optical intensity of sprite clusters when observed with TV-frame rate cameras is comparable with that of a moderately bright auroral arc, reaching 50 MR for the most intense events (Sentman et al. 1995).

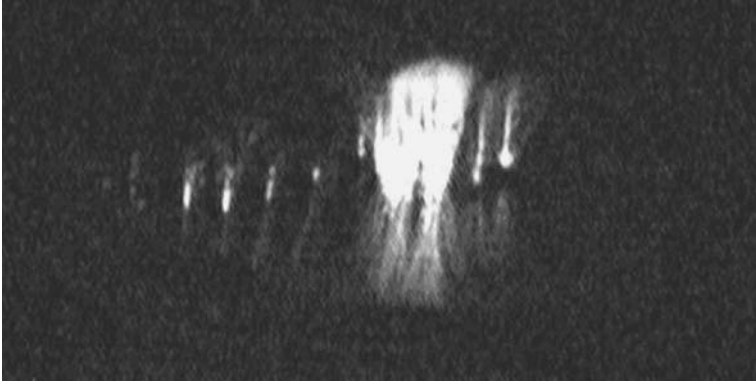
The structure of sprites is fascinating and is now being studied from the ground with increasingly sophisticated instrumentation. Early observations showed that the spatial structure could range from small single or multiple vertically oriented filaments of  $\sim 1$  km thickness, the c-sprites, to bright clusters of two, three, or more “carrot-shaped” emissions. Later telescopic imaging of sprites has revealed an amazing richness, with intertwined discharge channels and beads to scales down to 80 m diameter and smaller (Gerken et al. 2000; Mende et al. 2002).

High-speed photometry shows that the duration of sprites is from a few ms to  $\sim 200$  ms (Winckler et al. 1996; Armstrong et al. 1998), and high-speed imaging at 1 kHz reveals discharges, initiated at  $\sim 65$  km altitude, propagating downwards and, shortly after, upwards (Stanley et al. 1999; Stenbaek-Nielsen et al. 2000; Moudry et al. 2003). Even a frame rate of 1 kHz is not sufficient to capture fully the dynamics of sprites. Recent imaging at 10 kHz provides fascinating insights into the formation and propagation of streamers and resolves, for the first time, the streamer heads (McHarg et al. 2007). On the shortest time scales, sprite emission rates can reach 1–500 GR (Stenbaek-Nielsen et al. 2007).

Sprite morphology and temporal characteristics are certain to depend on the underlying thunderstorm source field; however, studies of this relationship are still in their infancy. Assumptions on the source field include the quasi-electrostatic field from a +CG discharge (Pasko et al. 1996, 1997), the field from the horizontal component of the discharge within the cloud feeding the +CG (Valdivia et al. 1998), and the very rapidly varying field composed of the directly emitted electromagnetic field from the intracloud (IC) flash current and the components reflected from the ground and the ionosphere (Cho and Rycroft 2001). One of the reasons that the source field is still debated is, of course, the difficulty of making in situ observations of the fields in the mesosphere.

It is also thought that local perturbations of the mesosphere may affect the sprite morphology and that understanding the sprite process will increase our understanding of the mesosphere itself. For instance, perturbations in the neutral gas density caused by gravity waves generated in the underlying thunderstorm, typically at scales of 10–50 km, could modulate the local threshold electric field for the sprite discharge, with low densities having lower thresholds (Sentman et al. 2003). An example of a sprite cluster that could have been formed in such a process is shown in Fig. 2. It is structured in individual sprite elements at horizontal distances of  $\sim 10$  km within an overall envelope of dimension  $\sim 100$  km that maps the region of the largest “reduced” field (namely, the electric field divided by the pressure) expected from a quasi-electrostatic field following a positive cloud-to-ground discharge (Neubert et al. 2005). However, it is not known whether the medium-scale structure (10 km) is caused by the local properties of the mesosphere or by interference patterns in the radiated and reflected source electric field as mentioned earlier.

It has also been proposed that perturbations of the electric conductivity induced by meteor trails at scales below 1 km are instrumental in the sprite ignition process (Suszcynsky et al. 1999; Symbalisty et al. 2000). Thus, the discussion of the medium and small scale structure of sprites and of the conditions that lead to ignition is relevant not only for



**Fig. 2** Medium-scale structure of sprites observed over Southern France from Observatoire Midi-Pyrénées on 28 August 2003 at 23:16:55.809 UT. The horizontal dimension of the luminous region is  $\sim 100$  km and the separation between sprite elements is  $\sim 10$  km (Neubert et al. 2005)

understanding the sprite discharge itself but also for advancing our understanding of the medium and small scale structure of the mesosphere.

### 1.3 Red Sprites: Discharge Mechanisms

Several theories have been put forward to explain the sprite phenomenon. They all rely on free electrons accelerated to energies sufficient to excite atmospheric constituents by collisions. However, they differ on the processes involved. They can be categorized according to the energy  $E$  of the free electrons *driving* the process observed as optical emissions:

- (1)  $E < \sim 15$  eV: optical emissions are really not signatures of a discharge but only of excitation of atmospheric constituents by collisions with free electrons created by detachment from negative ions.
- (2)  $\sim 15$  eV  $< E < \sim 100$  eV: free electrons are created by ionization of the neutral atmospheric constituents in a classical neutral air breakdown cascade.
- (3)  $\sim 5$  keV  $< E < \sim 5$  MeV: energetic seed electrons are created by cosmic ray ionization of the atmosphere. The gas breakdown is carried by electrons in this high-energy range.

The spectrum of sprites contains the information needed for studying the energetics of TLEs. Early observations suggested a strong signature of  $N_2(1P)$  emissions and an absence of the  $N_2^+$  Meinel emissions, consistent with average energies of free electrons of the order of 1 eV and insignificant amounts of air ionization. A theory has been proposed that invokes the release of electrons from pre-existing negative ions naturally present in the atmosphere (Swenson and Rairden 1998). The collisional energy needed for the release is 0.5–1.5 eV. This model belongs to category (1). It has been noted, however, that emissions in the Meinel bands, which reflect ionization processes, are strongly quenched by the atmosphere. Thus one has to be careful when comparing emission ratios estimated from ground measurements with theoretical predictions (Armstrong et al. 1998).

Models have been suggested where classical air breakdown is generated by the polarization electric field over the cloud tops following a positive cloud-to-ground stroke. The

cloud top charges to positive potentials, with time-scales of the order of hundreds of seconds, and the more tenuous atmosphere above has time to supply polarization currents to a shielding space charge above the cloud. When the positive cloud top layer discharges directly to the ground, a downward-directed, transient electric field is generated in the atmosphere above with a time constant that depends on the continuing current of the order of a few ms to a few tens of ms. Theory for conventional air breakdown requires 200–500 C to be discharged to ground in order for the field to be above the breakdown limit (Pasko et al. 1995, 1996, 1997). Breakdown occurs when streamers are formed with high space-charge density in the streamer tips. The downward-directed streamer propagating in the direction of the electric field is called a positive streamer. Positive streamers are thought to occur in the sprite discharge. The energy of electrons is generally higher in the streamer tip than in the streamer body (a sprite filament), and the tail of the electron distribution at energies exceeding the ionization threshold thus drives the avalanche. To explore the validity of the model, a characteristic energy of the electrons has been estimated and compared to estimates based on optical spectral measurements. Model and observations compare well with energies of  $\sim 2$  eV in the sprite proper (Morrill et al. 2002) and  $\sim 5$ – $25$  eV in the halo surrounding the sprite (Miyasato et al. 2003). The model is one of category (2) and is the one mostly accepted. It is worth commenting that the model also allows for discharges by electric fields of opposite polarity, i.e. directed upwards and generated by  $-CGs$ , although they would be less efficient than  $+CGs$ . However, negative  $-CGs$  have been observed to cause sprites on rare occasions (Taylor et al. 2008).

The runaway electron discharge process belongs to category (3). When the electron energy increases beyond  $\sim 500$  eV, the combined effects of the atmospheric collisional cross sections decrease and reach a minimum around  $\sim 1$  MeV. At higher energies the cross sections increase again because of the increase in the mass of a relativistic electron. The discharge process is driven by energetic electrons ( $\sim 1$  MeV) where collision cross-sections minimize for atmospheric constituents (Gurevich et al. 1992, 1999; Roussel-Dupré et al. 1994). In this energy range, the air breakdown field is a factor of 10 below the conventional air breakdown field. Because the threshold field is lower, this mechanism should be triggered before the conventional one described above, and should supply neutralizing currents keeping the field just at the relativistic breakdown limit. The runaway discharge process thus seems to preclude classical air breakdown (Roussel-Dupré and Gurevich 1996; Taranenko and Roussel-Dupré 1996). However, this scenario has now been revised. Simulations suggest two categories dictated by the duration of the parent lightning. For fast ( $< a$  few ms), high current ( $> 100$  kA) positive CG lightning, the high altitude discharge is dominated by a diffuse airglow driven by conventional breakdown at high altitudes ( $> 60$ – $70$  km), followed by runaway breakdown that initiates at the top of the cloud ( $\sim 15$  km) and develops upward to altitudes exceeding 90 km. For longer duration continuing currents ( $> a$  few ms), streamers driven by conventional breakdown develop at high altitudes ( $\sim 75$  km) and evolve downward to lower altitudes (terminating at  $\sim 50$  km). The streamers are then followed by a runaway discharge that is initiated at the top of the cloud and propagates to high altitudes (Roussel-Dupré 2005).

Observations indicate that several of the processes outlined above operate in sprite discharges, but their relative roles and importance remain unclear. One of the key parameters to be observed in order to evaluate the significance of the processes is the degree of ionization. Models belonging to category (2) are in fairly good agreement with spectral observations of  $N_2^+$  ionization transitions. However, such observations are



marginal from the ground or from aircraft because of atmospheric absorption at the blue end of the spectrum. Spectral observations are best performed from space. That has been done since 2004 by the Instrument for Sprite and Upper Atmospheric Light (ISUAL) instrument on the FORMOSAT-2 satellite (Mende et al. 2005; Frey et al. 2005), which clearly show the existence of significant impact ionization of  $N_2$ . Ionization has also been inferred from ELF observations of currents associated with sprites (Cummer and Inan 1997; Cummer 2003), from radar echoes from ionization enhancements thought to be generated by sprites (Roussel-Dupré 1997), and from perturbations of VLF transmitter signals propagating in the Earth-ionosphere waveguide (Rodger 1999; Haldoupis et al. 2004). However, while classical electric breakdown of the atmosphere has been demonstrated beyond doubt, the role of the release of low energy electrons from negative ions in the sprite process is not known.

It is difficult to make progress on the question of the importance of the runaway process of category (3) without observations from space of bremsstrahlung radiation in the upper atmosphere, as the radiation is absorbed in the lower atmosphere and is not observable from the ground. The observations must be taken simultaneously with observations of optical emissions to identify the connection of TGFs to lightning and TLE activity. Such observations have so far not been performed, but are planned with the French microsatellite TARANIS (Tool for the Analysis of RAdiations from lightNings and Sprites) and the ASIM (Atmosphere-Space Interactions Monitor) payload on the International Space Station, both scheduled for launch in 2012. In the mean time, experimental evidence for this process operating in relation to lightning is mounting, with observations of X- and  $\gamma$ -rays associated with thunderstorm electrification and lightning discharges from balloons, from mountain tops, and in connection with rocket-induced lightning (Brunetti et al. 2000; Moore et al. 2001; Dwyer et al. 2004a, b, 2005).

The runaway process requires the presence of seed electrons at energies within the runaway regime and with their momentum vector within a certain angle of the downward-directed electric field. Such electrons are abundant from cosmic ray bombardment of the atmosphere. Computer simulations of sprites indicate that  $10^6$ – $10^7$  electrons/cm<sup>2</sup>, accelerated in the  $\sim 100$  MV potential between cloud tops and the ionosphere following a lightning discharge, are injected into the magnetosphere (Lehtinen et al. 1996, 1997). One of the questions raised by the RHESSI observations is the lack of TGF observations over North America, where powerful thunderstorms develop in the summer time and sprites are observed routinely. This is particularly puzzling when one considers the effect of the Earth's magnetic field, which is horizontal at the magnetic equator, and has increasing inclination towards higher latitudes. For relativistic electrons with energies around  $\sim 1$  MeV, the geomagnetic field in the equatorial region begins to play a role at altitudes of  $\sim 30$  km. Above this altitude the almost horizontal geomagnetic field hinders a runaway electron discharge driven by a vertical electric field. Thus, the runaway mechanism favours regions of high inclination of the geomagnetic field because it allows electrons to be accelerated upwards. To resolve this issue an alternative acceleration process to the quasi-static electric field has been proposed. It invokes the field associated with the return stroke of the lightning process. This field has a horizontal component, and thus a component parallel to the geomagnetic field at low geomagnetic latitudes (Inan and Lehtinen 2005). However, it is still not certain at what altitudes TGFs are generated. Some observations suggest altitudes below 30 km within, or just above, the clouds although most of the radiation then is absorbed in the atmosphere before reaching satellite instrumentation (Østgaard et al. 2008). In this case the Earth's magnetic field would play no significant role in the discharge process.

## 1.4 The Eurosprite Campaigns

While observations of sprites made in the 1990s concentrated on the very powerful thunderstorms over the plains of the USA, it was not known whether smaller storm systems like those in Europe would generate Transient Luminous Emissions (TLEs). To address this question, the first ground campaign in Europe was conducted during the summer of 2000. The campaign documented sprites over Southern France and Northern Italy and the observations suggested that sprites were more common than previously thought (Neubert et al. 2001).

Since then, annual observational campaigns have been conducted over Southern Europe every summer. The backbone instrument is a camera system mounted at the Observatoire Midi-Pyrénées, in France. It has been upgraded every year with varying combinations of cameras and photometers. The cameras operate at 40 frames/s with light-sensitive CCDs at up to  $1000 \times 1000$  pixels. The system is connected to the Internet and commanded from anywhere with access to the Internet. Real time information on lightning activity is available from the French Météorage company, allowing the system managers to point the cameras in the direction of lightning activity. Compressed event images are available in real time, with uncompressed data and extended footage downloaded during the daytime.

With the European Framework 5 Research Training Network “Coupling of Atmospheric Layers” (CAL) (2002–2006), the campaigns gained considerable momentum. It brought together teams from different scientific disciplines with expertise in a wide range of experimental techniques. The guiding principle for the network campaigns was that new science comes from new observations or from new combinations of observations. Following this approach, the *EuroSprite* campaigns included infrasound measurements from several locations in France, documenting the infrasound signatures of sprites at up to 1000 km distance, and measurements of ground-based VLF (3–30 kHz) transmitter signals by a receiver on Crete, showing impulsive perturbations in the amplitude when the signals propagate in the Earth-ionosphere waveguide over thunderstorms. In addition, measurements of electromagnetic waves from a number of locations ranging from ULF (<3 Hz) to VHF (30–300 MHz) have given information on lightning activity, such as charge moment changes, intracloud lightning activity, and perturbations to the ionospheric conductivity. The collaboration also enabled more optical stations to be set up, which increased the coverage and in some cases permitted the triangulation of sprites when observed from two stations. Finally, meteorological measurements made from satellites or by ground-based radars of cloud cover, cloud top temperatures, and precipitation have also been available.

Sprite observations and associated lightning activity during the summers of 2003–2006 form the basis of the experimental observations presented in this paper. The data are summarized in Appendices 1–3. The first results of the *EuroSprite2003* campaign were published by Neubert et al. (2005). In the following sections we report on the advances made since then, both as a result of campaign activities and modelling efforts by participants in the network.

## 2 The Thunderstorm Source of Sprites

Operational lightning detection systems usually measure the properties of cloud-to-ground (CG) lightning, whereas intracloud (IC) lightning activity remains undetected. For this reason, studies in the past of the relationship between lightning discharges and sprites have focussed primarily on the role of CGs, showing that, in almost all cases, a +CG can be

identified as a trigger of sprites (Boccippio et al. 1995). Broad-band recording of radiation at frequencies below  $\sim 100$  Hz in the ELF range, which can be observed at very considerable distances from the source, have further enabled estimation of the associated charge moment change (CM), which is the amount of charge lowered to the ground multiplied by the distance between the charge reservoir and the ground. It has been shown that the CM change needed to generate sprites is of the order of 300–600 C km for short delayed sprites (ms) and larger for larger delays (Cummer 2003) (Note that some authors consider the charge moment change to be twice this, accounting for the image charge in the ground—see also Sect. 7.2).

It was recognized early (Lyons 1994; Sentman et al. 1995) that sprites are associated with the stratiform regions of mesoscale convective systems adjacent to the stronger convective cores and that they tend to occur during the mature to decaying stages of thunderstorm development. These regions contain extensive layers of charge that can be tapped by IC lightning. Rare observations from IC lightning mapping systems have shown that a significant IC component may be feeding the +CG flashes that trigger sprites and thus sustaining strong continuing currents (Stanley et al. 2000; Lyons et al. 2003).

During the *EuroSprite* campaigns thunderstorms were monitored with a more complete set of instrumentation, thereby allowing a better characterization of cloud and lightning properties of the sprite-producing regions of the storms. The storms were sometimes located over the short-range lightning detection system SAFIR that also detected IC discharges. The SAFIR systems were located in northern Spain and in the Provence region of Southern France. The systems were based on three VHF ( $\sim 100$  MHz) receivers in a triangle with  $\sim 100$  km distance between the receivers that measure activity in clouds, giving information on the location (in the horizontal plane) and time of IC and CG activity (Drüe et al. 2007). Additional information came from a broad-band, VLF ( $\sim 10$  kHz) receiver located at Nançay in central France. CG activity was also available from conventional lightning detection networks provided by the French Météorage system of 17 sensors covering the whole country. Each sensor uses two techniques for CG location, a direction finder (DF) and time of arrival (TOA) information. The characteristics of the CG flashes determined include the location (horizontal), the time of occurrence, the polarity, the multiplicity and the peak currents of each stroke (Cummins et al. 1998).

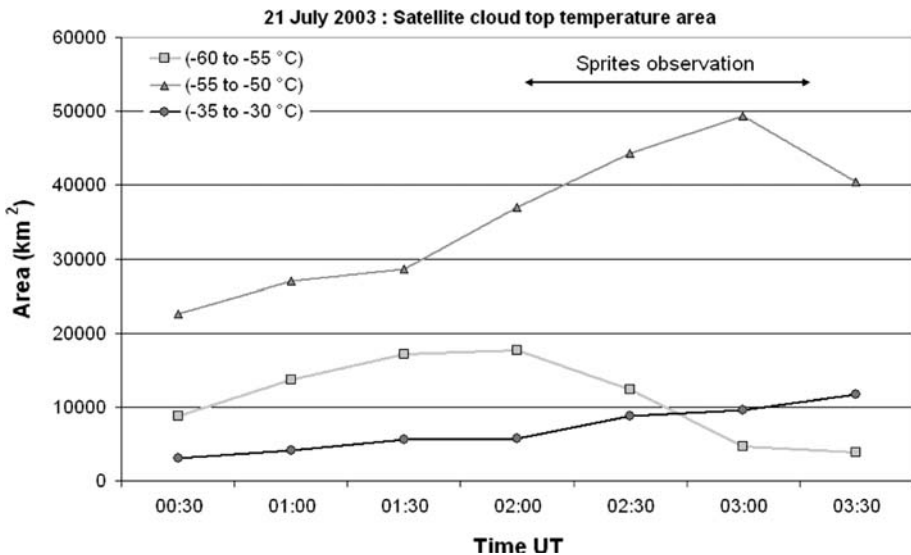
Other characteristics of thunderstorms are derived from meteorological radars. France is covered by a network of weather radars, ARAMIS, run by MétéoFrance. Seventeen radars with 250-km range provide complete horizontal scans every 5 minutes. The elevations of the beams are kept almost constant so that altitude information is not available. The parameter provided by the radars is the reflectivity factor which is directly related to the precipitation rate (or rainfall rate). Sprites can then be related to the cloud system properties, to determine the type of thundercloud structure that favours their production and in what evolutionary phase of the convective system they are generated. Near-infrared images from the geostationary Meteosat satellite are also used, from which the altitudes of the cloud tops are estimated.

In the following sections we give results from our studies of the characteristics of sprite-producing thunderstorm regions observed in Europe and the importance of the IC activity for sprite generation and morphology. Some of the results are published in van der Velde et al. (2006) and Soula et al. 2008 whereas others are published here for the first time. Earlier studies, including those made by (Lyons 1996) and Lyons et al. 2003, analyzed storms over the US Plains in depth.

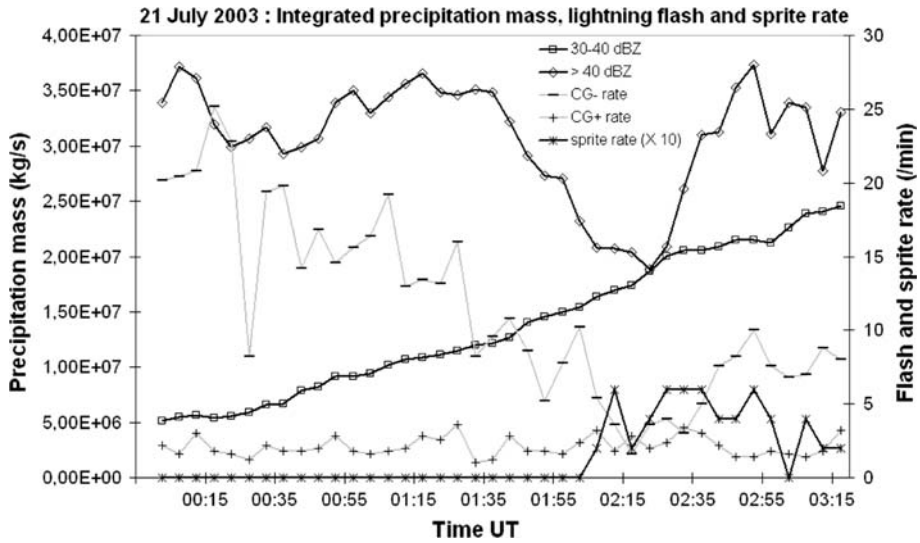
## 2.1 Characteristics of Sprite-Producing Storms

Optical imaging, which is the most reliable method of sprite detection, can only be done during night time when the sky is clear and there is an unobstructed view of the region observed. One usually cannot be certain that the full period of sprite production by a storm is recorded, as sprites could be undetected because they were outside the field of view of the cameras. While the complete thunderstorm phase of sprite production from start to end usually cannot be characterized with certainty, it is possible to determine the thunderstorm conditions present during the period when sprites are observed. Cases were selected for studies based on the number of sprites observed in a given period and the availability of simultaneous radar, satellite and lightning data. The last requirement excludes most of the cases over the Mediterranean Sea. In all, seven of the storm systems observed during 2003, 2005 and 2006 were analysed with regards to cloud properties and their relation to sprite generation (Soula et al. 2008).

Satellite images suggest that sprites tend to occur over the stratiform region near the coldest (highest) cloud tops. For the storm systems observed during the *EuroSprite* campaigns, sprites tended to occur when the cloud canopy reached its largest extent. The size of the storm with cloud top temperatures lower than  $-50^{\circ}\text{C}$  (near the tropopause level) peaks during the sprite period. The significant growth of this region may be attributed to the development of a front-to-rear mesoscale updraft. Figure 3 shows the development of cloud top regions at various temperatures during a sprite-active phase of the thunderstorm of 21 July 2003. The maximum size of the  $-50^{\circ}\text{C}$  region lags that of the coldest cloud tops (over the strongest cells, usually  $-60^{\circ}\text{C}$ , or colder) by an hour or more. A similar study of a mesoscale convective system over Kansas found the maximum sprite production rate when the overlying cloud top temperatures were minimum,  $-69^{\circ}$  to  $-72^{\circ}\text{C}$  (São Sabbas



**Fig. 3** Time series of the area of cloud regions with different cloud top temperatures derived from observations made by Meteosat for the storm on 21 July 2003. Sprites are generated in the period from  $\sim 2:00$ – $3:30$  UT when the area of the coldest cloud tops decreases and the area of medium-cold cloud tops reaches its maximum



**Fig. 4** Time series of the integrated precipitation mass from regions of the cloud with different radar reflectivity values, of the positive and negative cloud-to-ground lightning flash rates, and of the sprite rate for the storm on 21 July 2003

and Sentman 2003). The difference is possibly caused by a difference of tropopause height for the two regions or by differences in storm size.

Radar reflectivity  $Z$  (expressed in dBZ) gives a view of the internal structure of thunderstorms. Reflectivity  $Z_0$  can be conveniently converted to the rainfall rate  $R$  by the commonly used empirical  $Z$ - $R$  relationship stemming from the Marshall and Palmer 1948 drop size distribution:  $Z_0 = 200 R^{1.6}$  ( $Z_0$  in  $\text{mm}^6 \text{m}^{-3}$  and  $R$  in  $\text{mm/hr}$ ). The rainfall rate can be summed over an area and be expressed in  $\text{kgs}^{-1}$ . This was done for different intervals of  $Z$  values in order to separate convective and stratiform rainfalls. The result for the storm of 21 July 2003 is shown in Fig. 4. All storms begin small with predominantly intense convective precipitation ( $>50$  dBZ). Sprites were observed during periods where the area of the stratiform precipitation region with reflectivities from 30–45 dBZ was growing, and usually lasted until the size of this region began to decrease. This result is consistent with the observed range of reflectivity values for sprite-triggering +CG flashes over the US Plains reported by Lyons et al. 2006. A decrease of the area suggests that the regeneration of precipitation in the stratiform region, and hence of fresh charge, has ceased. Areas of reflectivity weaker than 30 dBZ were not found to correlate with sprite production in any of the observed systems.

The precise reflectivity range that correlates with sprites in the manner described above differs among the storms. The storms of 21 July 2003, 29 July 2005 and 11 September 2006 featured the strongest growth in the 30–40 dBZ range, during which time sprites were observed. The area of the 40–45 dBZ region was growing, especially in the late stages of the storm, but not correlating with the sprite period. On the other hand, the storm of 17 November 2006 already had a very extensive area of 30–40 dBZ reflectivity when sprites were first observed, and the 40–50 dBZ-range expanded its area significantly during the sprite period; sprites stopped when this area started to decrease. This relatively strong stratiform precipitation was enhanced by orographic lifting of moist Mediterranean air over

the Massif Central plateau. However, the +CGs that triggered sprites in this system occurred in the usual 30–40 dBZ range, bordering the higher reflectivities.

For the storm of 21 July 2003, the start of the sprite period could unambiguously be identified from optical observations. It began not long after a part of the convective region weakened, as made evident by the rapid decrease of >45 dBZ integrated precipitation. The lightning flash activity of the storm is shown in Fig. 4. Weakening of the convective region is accompanied by decreasing –CG flash rates, to very low levels, when sprite-triggering +CGs started. During the sprite period, low flash rates of predominantly positive flashes are observed. This is a trend also observed in other storms. This observation is different from the large MCS studied by São Sabbas and Sentman (2003) where sprites were found to peak during the period with the largest –CG flash rates. However, in such a large MCS there can be several convective core areas in different stages of their life cycle. In the 21 July storm convective cells were regenerating themselves in an area far from where the sprites occurred, so the rising –CG rate shown on the figure is not relevant to the sprite producing region.

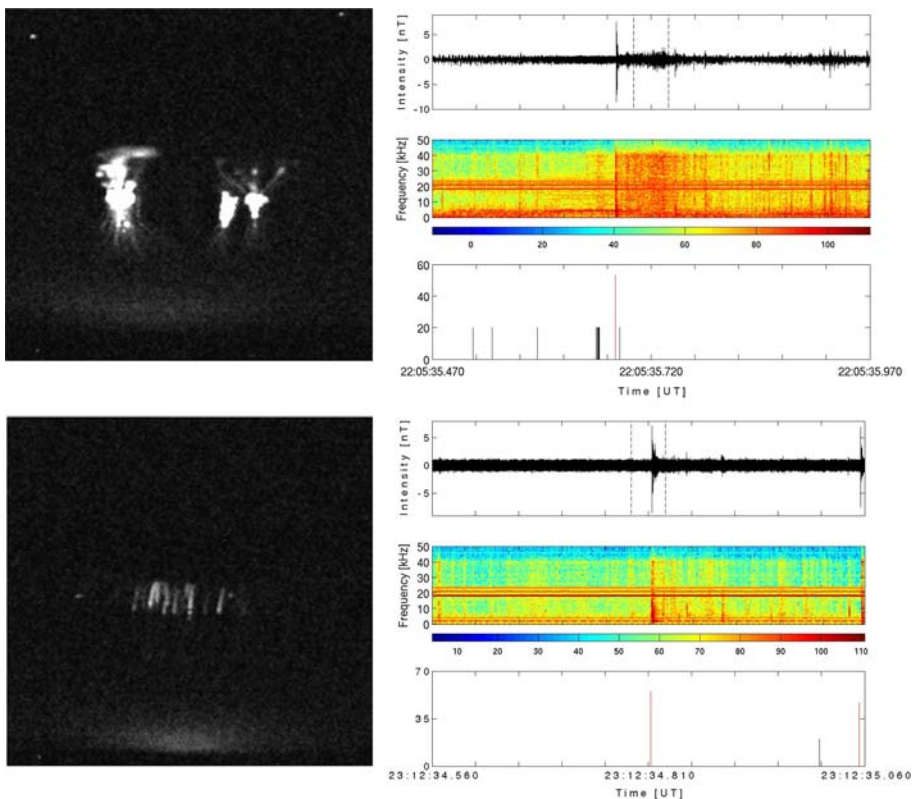
A statistical analysis of all CG lightning flashes in a thunderstorm has been undertaken for seven sprite-producing storm systems. Within circles of 5 and 10 km radii, radar reflectivity and cloud top temperatures were determined for every CG flash. An average precipitation rate over the circle was determined using a Marshall–Palmer relation for the raw reflectivity average ( $Z$ ). Radar pixels with a reflectivity lower than 9 dBZ ('dry') in the circle have been excluded. Next, flashes were grouped into three basic classes: flashes not related to sprites (+, –), flashes occurring within 250 ms before a sprite assumed to be the triggering flash (ST+), and extra flashes that appeared close to the sprite in time and space (S+, S–). This latter category occurred usually within 1 second before or after the sprite and in areas with sparse lightning activity, giving the appearance that they were part of the same horizontal lightning flash sequence. The results show that ST+ and S+ occur in the most uniform precipitation region of moderate intensity, whereas (+) occur in a wider variety of conditions but overlapping with those associated to sprites. One-third of +CG flash sequences were found to produce two-thirds of all sprites. Of the sequences, first +CGs occurred in less stratiform precipitation regions than later +CGs in the sequence, suggesting that extensive horizontal lightning was associated with the +CGs, propagating away from precipitation cores. This interpretation is supported by earlier observations of the evolution of CG and IC lightning channels (Carey et al. 2005) and IC and +CG activity in association with sprites (Lyons et al. 2003).

## 2.2 Cloud Discharges and Sprites

The relationship between sprites and their causative discharges has been studied using the entire data set of more than several hundred sprites. The distribution of the time delays between sprites and their causative +CGs was found to have a well defined tail showing that at least one third of the sprites were lagging the causative +CG discharges by more than 25 ms, reaching values up to ~250 ms (Mika et al. 2005). Based on observations of ELF/VLF observations it has been suggested that IC activity plays a role in sprite generation (Ohkubo et al. 2005). This was examined for 15 sprites within range of the SAFIR VHF lightning detection system which directly could monitor IC activity. Column sprites were generally characterized by a short time delay relative to the causative +CG discharge (<25 ms), high +CG peak current, little IC activity and short, intense bursts of broad-band VLF radio wave activity. Carrot sprites seem to be associated with longer time delays, large IC activity and weaker, but longer-lasting, bursts of lightning discharges (termed

atmospherics, or sferics for short) at times preceding the +CG (van der Velde et al. 2006). The association of lightning activity, sferic activity and sprites is illustrated for two cases in Fig. 5. The above characteristics suggest that IC discharges and long-lasting continuing currents play an important role in the generation of carrot sprites but are relatively unimportant for the more impulsive column sprites. These characteristics are summarized in Table 1.

The IC flash and sferic activity is caused by breakdown processes inside the clouds feeding the continuing current of the +CG. As the sferic radiation is broad-band, this may not seem to favour the mechanism proposed by Cho and Rycroft (2001), where ionization enhancements in the lower ionosphere are caused by structured interference patterns of fields from horizontal lightning bouncing in the Earth-ionosphere cavity. However, a time-varying interference pattern will be created by the radiated waveform, whatever its



**Fig. 5** Observations made on 23 July 2003 (van der Velde et al. 2006). Magnetic East-West component of ELF/VLF waveforms (top panel, right), corresponding wave spectra (middle panel, colour) and lightning activity (bottom, right) for 500 ms time period around observations of carrot sprites (top, left) and column sprites (bottom, left). The two vertical dashed lines on the waveform panels mark the time of sprite detection by the optical cameras. The ELF/VLF data were observed at Nançay, France, about 300 km from the events. The lightning data of +CG strokes from Météorage are shown in red and the SAFIR VHF intracloud lightning sources in black. The vertical scale indicates the +CG peak current (kA); the lengths of the black lines representing the SAFIR data are scaled arbitrarily. For the carrot sprite significant broad-band ELF/VLF activity is following the causative +CG and some SAFIR activity is preceding it. For the column sprites the broad-band activity is weaker and SAFIR sources are absent

**Table 1** Characteristics of the lightning source of column and carrot sprites

Type	Time delay	VHF Sources SAFIR	VLF sources Nancy	Peak current	Triggering lightning
Column	Small (<25 ms)	Low activity	Short (<25 ms) high intensity	High	Impulsive +CG
Carrot	Larger	High activity preceding +CG	Long duration lower amplitude	Moderate	+CG and IC

bandwidth, so that this mechanism should, nonetheless, operate. One observed sprite event could not be associated with a +CG flash, but was instead preceded by an unusually large cluster of SAFIR IC sources and an unusually strong VLF sferic cluster, suggesting that IC lightning activity can also be important for the triggering of sprites.

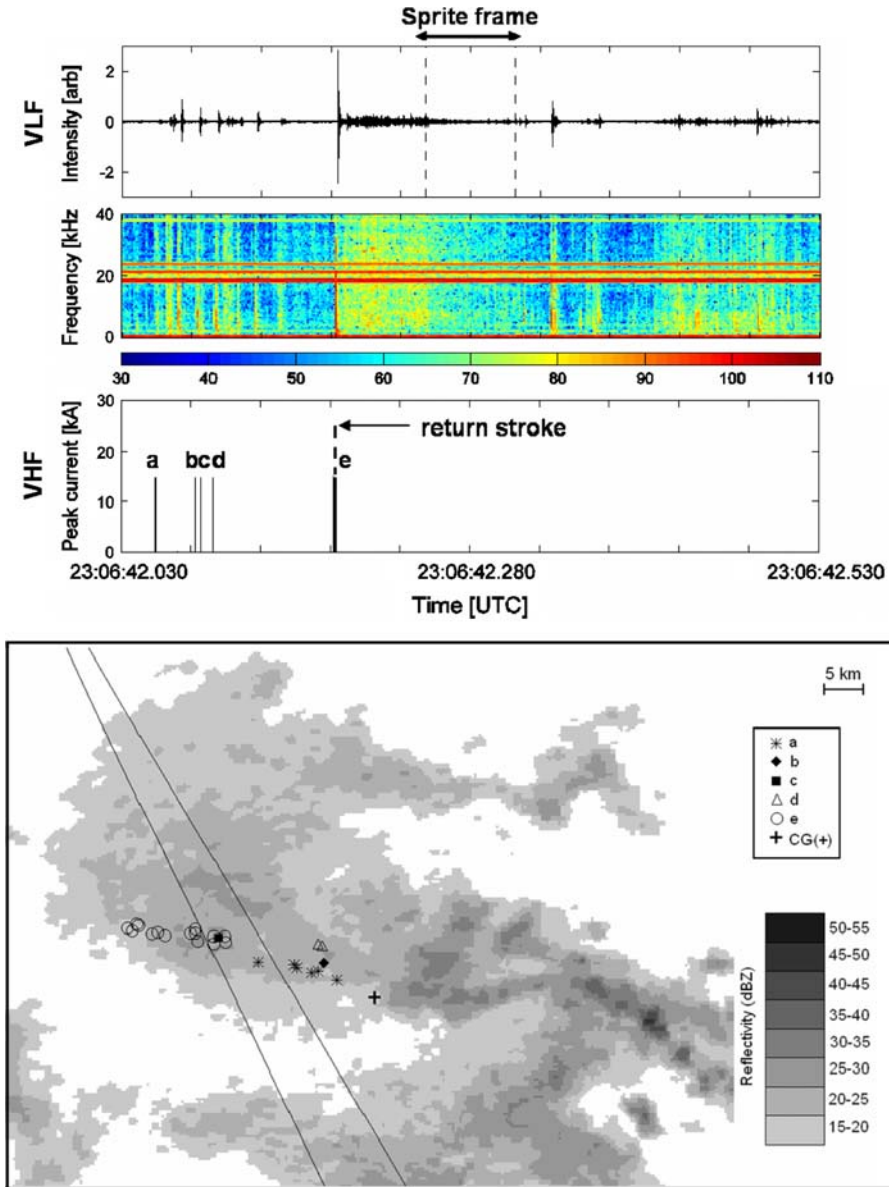
An alternative interpretation is that a +CG flash was indeed generated, but went undetected, which opens the question of the reliability of lightning detection networks. The VHF sources observed by the SAFIR system were usually found to precede the +CGs, but were never observed after the +CG stroke. In contrast, the VLF broad-band sferic activity usually began at the onset of the +CG flash. This suggests that the interferometry system cannot locate sources when the fields become too complex, with high currents established by the +CG and new IC discharges being formed simultaneously at widely different locations in the storm. This conclusion is in accordance with (Mazur et al. 1997, 1998).

The relationship between IC lightning and the horizontal displacements of a sprite from its parent +CG stroke was studied for a sprite displaced by 20 km which was observed over northeastern Spain on 16 August 2005. This was within range of one of the SAFIR systems and the relevant information is shown in Fig. 6. First the VHF sources propagated from the location of the (later) +CG towards the region over which the sprite would occur, with a speed that is typical of a horizontal stepped leader ( $2 \times 10^5$  m/s). The +CG flash was then triggered and was followed by broadband VLF activity and the sprite. The sequence suggests that IC lightning propagated into the stratiform precipitation region as a negative stepped leader, which connected to the ground by the +CG discharge near the location where the negative leader started. A similar observation was reported by Mazur et al. (1998). The return stroke process discharges the IC channels, bringing their potential closer to that of the ground, and triggering additional breakdown that supplies charge to the +CG channel over a time interval as long as the sferic cluster (60 ms). This event shows the well known fact that a +CG may transfer charge from a remote region of the storm to the Earth. It also shows that the process may favour sprite generation which is laterally displaced from the +CG. Our data suggest that this is a very common occurrence. The above interpretation is consistent with the conclusions of Marshall et al. 2007 that VLF sferic clusters are produced by horizontally propagating lightning, although their data do not show a one-to-one match between all lightning data detected by the Lightning Mapping Array and the VLF sferic clusters.

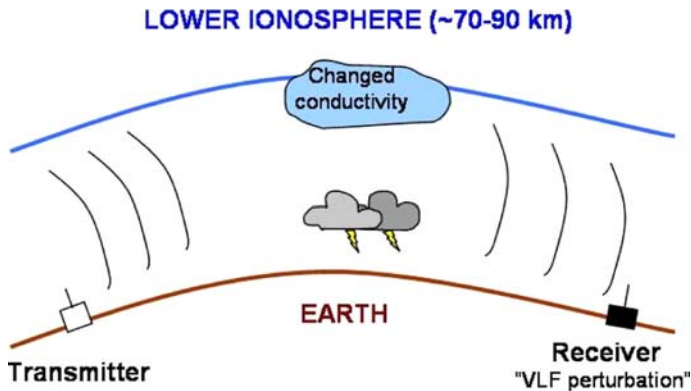
### 3 Ionization and Heating of the Mesosphere by Sprites, Elves and Lightning

Studies of conductivity changes above thunderstorms in the mesosphere and lower ionosphere were undertaken by measuring changes in the amplitudes of signals from narrowband radio transmitters passing over thunderstorm regions. The measuring principle





**Fig. 6** Observations on 16 August 2006: Upper panel top and middle: Broadband VLF waves (East-West component) 250 ms before and after the sprite event at 23:06:42.260 UTC. The two vertical dashed lines indicate the sprite observation time, accounting for the video frame integration time of 40 ms plus the timing uncertainty of 12 ms. Upper panel bottom: Cloud-to-ground strokes detected by Météorage (red) and intracloud lightning sources detected by SAFIR (black). The vertical scale indicates the +CG peak current (kA). The black lines representing the SAFIR data are scaled arbitrarily; a, b, c, d and e mark the different sequences of the activity in VHF range. Lower panel: Lightning activity and composite radar reflectivity distribution. The small symbols indicate the VHF sources gathered in five sequences (a, b, c, d and e of upper panel) and the arrows indicate the direction of displacement. The cross indicates the location of the +CG stroke. The two lines show the directions of the sprite edges as seen from the Pic du Midi



**Fig. 7** Diagram showing how the bottom side ionosphere and the mesosphere above a thunderstorm are probed by the signals from a ground based radio transmitter propagating in the Earth-ionosphere waveguide

is shown in Fig. 7. The transmitters were either powerful VLF (3–30 kHz) transmitters used for navigational purposes or MF (0.3–3 MHz) radio station transmitters. At these frequencies electromagnetic waves propagate in the Earth-ionosphere waveguide, the upper boundary of which is affected by perturbations of the electrical conductivity in the mesosphere where TLEs (i.e. sprites and elves) are generated. State-of-the-art VLF receivers in Crete and France recorded simultaneously with the TLE detection system at the Observatoire Midi Pyrénées, and with lightning detection systems; broad-band MF radio signals were recorded at several locations in France.

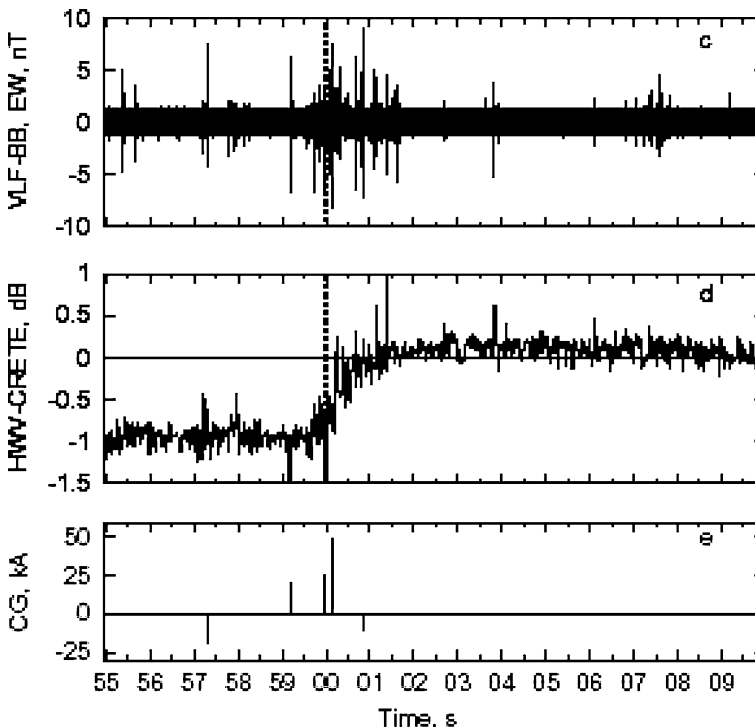
### 3.1 Perturbations by Sprites

The electromagnetic energy from a lightning discharge can cause additional ionization and heating of the electron population in the collisional plasma of the lower ionosphere. Consequently, the resulting change in electrical conductivity may cause changes in the amplitude and/or phase of VLF ground transmitter signals passing through the region and can be detected experimentally provided that the perturbation is of sufficient magnitude (Rodger 1999). The most common perturbation observed follows within a few ms of the causative lightning discharge and has a short onset duration of less than  $\sim 50$  ms. Such perturbations are referred to as “early/fast” (Inan et al. 1995). During the nights of 21–24 July 2003, 42 sprites were observed over three thunderstorms in central France. They were located such that VLF ground transmitter signal paths, received in Crete, passed over the thunderstorms. All these sprites were triggered by +CGs and the receiver on Crete detected “early” VLF amplitude perturbations for  $\sim 90\%$  of them (Haldoupis et al. 2004; Mika et al. 2005). On the other hand, among the much more numerous +CG and –CG discharges which were not associated with sprites, only a few were associated with amplitude perturbations. Video frames, available for one of these events, were played back; this led to the discovery of an additional sprite which had previously been missed. These observations suggested that sprites are nearly always accompanied by “early” VLF perturbations.

Most of the sprite related VLF events had long recovery times ( $\sim 30$ – $300$  s), which suggests spatially extended, diffuse regions of electron density increases at altitudes higher than 75 km. This agrees with previous studies and supports theoretical predictions of air

breakdown in the upper D-region ionosphere during sprite occurrences, triggered by strong quasi-electrostatic (QE) fields. For more details see Mika et al. (2005) and Sect. 7.2.

A new category of VLF perturbations with a long onset duration of up to 2.5 s was observed in relation to  $\sim 55\%$  of the sprite events. These “early/slow” events are indicative of a new physical process at work which, following a causative +CG discharge, leads to a slow change in the conductivity of the lower ionosphere. Analysis of simultaneous broadband VLF recordings from the receiver in Nançay located within a few hundred km of the storms shows that the growth phases of “early/slow” events coincide with the occurrence of continued lightning activity during the growth phase. An “early/slow” event seen by the Crete receiver is shown in Fig. 8 together with Nançay broadband time series and Météorage lightning data. The event is associated with a few sequential CG lightning strokes (lower panel), and clusters (bursts) of VLF sferics which are due to intracloud (IC) lightning (upper panel). The observations suggest that the long onset duration is a result of secondary ionization build-up in and/or heating of the upper D region below the nighttime VLF reflection height, caused by sequential electromagnetic pulses (EMP) radiated upwards from horizontal IC discharges. The radiation fields accelerate sprite-produced electrons which then impact to produce secondary electrons in an avalanche process which



**Fig. 8** An “early/slow” VLF event observed in Crete (middle panel) at  $\sim 2000$  km from the thunderstorm. A sprite is observed at the time marked by the vertical dashed line; currents flowing in the causative +CG discharge and other CGs are shown in the bottom panel. This “early/slow” event is characterized by an onset of  $\sim 2$  s, during which sferics (top panel) are detected by a broadband VLF receiver in France  $\sim 200$  km North East of the thunderstorm. The sferics are attributed to intracloud lightning discharges which heat and ionize the lower ionosphere, causing the observed onset in the signal from the HWV transmitter at Le Blanc (Rosnay), France, to Crete

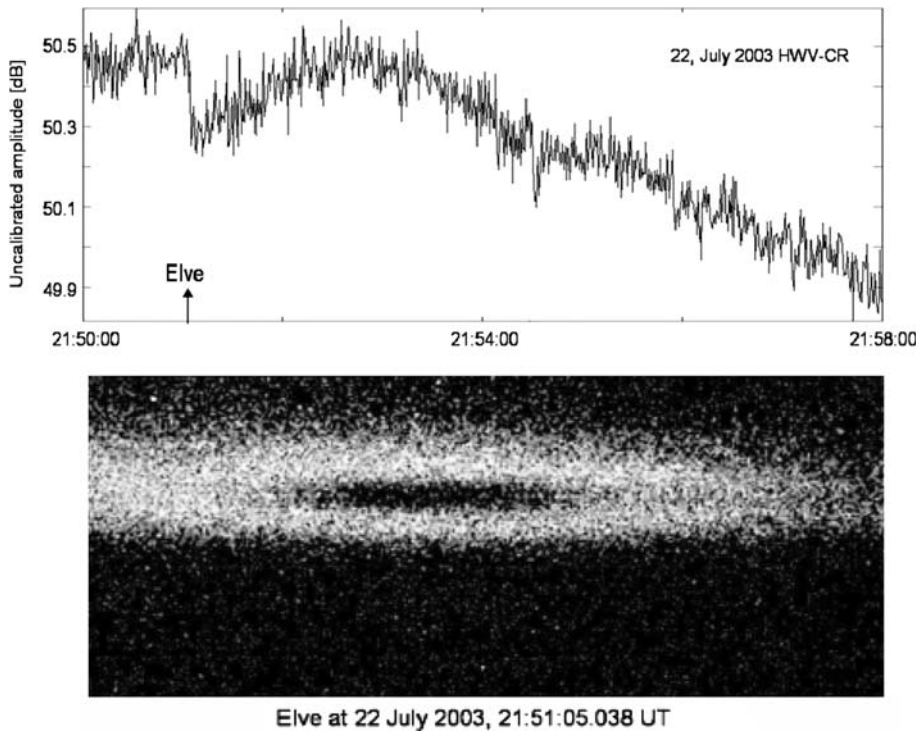
leads to the observed ionization build-up (Haldoupis et al. 2006). However, any associated optical emissions are below the sensitivity of the cameras (magnitude 6).

VLF transmitter signals recorded by the Nançay broadband VLF receiver, located 200 to 300 km West of the sprite-producing storms of 21 to 24 July 2003, were analyzed to study VLF backscatter from sprite-induced ionisation structures, as suggested by Dowden et al. 1996. Our study has shown that only ~5% of the sprites were accompanied by weak backscatter, indicating that backscatter from sprite plasma structures is possible but not as frequent as recorded by Dowden et al. (1996).

An existing D-region model, suitable for simulating electron density variations, was applied in order to compute the relaxation of extra ionization associated with sprites. The recoveries of observed “early” VLF perturbations were compared to the model results and the magnitude and altitude of enhanced ionisation was inferred. Reactions between four types of charged particles (electrons, negative and positive ions and positive cluster ions) were accounted for, along with various electron loss mechanisms. Direct heating effects were not included since “early” VLF perturbations are known to be caused by electron density changes (electron cooling times at the altitudes of interest (~75–90 km) are between ~5 ms and 0.2 s (Rodger et al. 2004), much shorter than the “early” VLF event recovery times). The continuity equations describing the time variation of the number densities of each of the species were solved numerically to compute the decay with time of the electron density enhancements produced by the action of lightning-induced QE fields which can lead to the appearance of sprites. The model results, obtained using different parameters corresponding to various altitudes, were then compared with the recovery times of “early” VLF events observed in relation to sprites. Good agreement with the observed recovery times of 20 to 250 s was obtained assuming electron density increases at altitudes of 75 to 85 km by factors from  $10^2$  to  $10^4$  times the unperturbed densities. This, in addition to the strong forward scattering pattern observed in relation to the VLF events recorded on Crete, leads to the conclusion that the incident VLF transmitter signals underwent scattering from horizontally extended diffuse regions of electron density enhancements, most likely associated with halos or the diffuse region of the upper part of carrot sprites, rather than small-scale streamers observed at lower altitudes (Mika 2007).

### 3.2 Perturbations by Elves

Elves are generated by the electromagnetic pulse (EMP) radiated by an intense CG stroke (e.g., Rowland 1998). It has been suggested that the EMP also creates ionisation (Tararenko et al. 1993; Rowland 1998), but this has not been verified experimentally. The largest database, so far, of elves observed over thunderstorms also probed by VLF transmitter signals, was secured during the *EuroSprite* campaigns. The optical data came from the cameras at the Observatoire Midi-Pyrénées (five elves) and from the ISUAL instrument on the FORMOSAT-2 satellite (Chern et al. 2003). VLF perturbation studies were conducted to search for evidence of ionization produced simultaneously with the optical emissions of the elves (Mika et al. 2006). During the *EuroSprite* 2003 campaign all five elves captured were found to be accompanied by “early” VLF perturbations on transmitter signals recorded by the Crete receiver. An example is shown in Fig. 9. Many elves observed from space, about 20 over Europe and 280 over North- and Middle-America, were examined. Eight elves detected above Europe and 42 seen over America were located at distances less than 500 km from the transmitter-receiver great circle paths, but only a few (all seen above Europe) were associated with early VLF perturbation events.



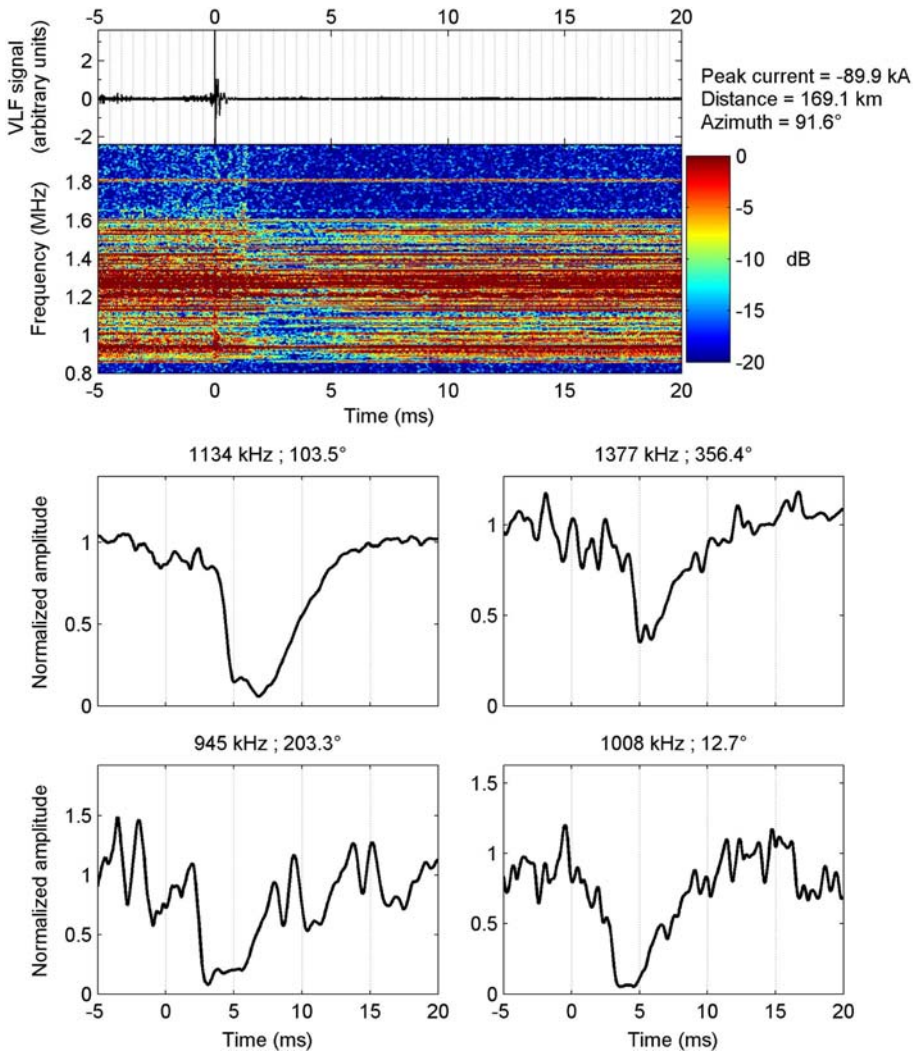
**Fig. 9** An “early/fast” VLF perturbation event (upper panel) detected by the Crete receiver  $\sim 2000$  km from the thunderstorm, and an associated elve observed from the Observatoire Midi-Pyrénées shown in the lower panel. This provides evidence for the creation of a diffuse region of heating/ionization in the upper D region which produces strong forward scattering of the incident VLF wave

The observations suggest that the lightning EMPs at times create regions of diffuse electron density enhancements in the uppermost D region of the ionosphere and that sub-ionospheric electron density changes associated with elves can intrude to lower altitudes (say below 85 km) and perturb VLF transmitter signals. VLF detection, however, is sensitive to several factors such as the distance of the elve to the receiver, the altitude of the perturbed ionospheric region and the EMP intensity. The EMP may be sufficient to cause optical emissions (elves) and ionization, but not ionization at lower altitudes which affects VLF propagation.

### 3.3 Perturbations by Lightning

Broad-band Medium Frequency (MF) measurements can also be used to study the perturbations of the lower ionosphere related with lightning activity. The Avril sur Loire (AVF) station measures simultaneously the vertical electric field in the VLF and MF band (500 kHz–1500 kHz) using an active rod antenna mounted on a 3-m-high mast. It was set up in the centre of France ( $3^{\circ}21.075$  E,  $46^{\circ}47.352$  N) in 2004. During the summer, no sprites were recorded; however, the effect of lightning from several storms was analysed.

It was very frequently found that, just after a return stroke, strong attenuation of narrow-band, MF radio signals was observed, as shown in Fig. 10. This perturbation effect has



**Fig. 10** The top of the plot is the VLF signal recorded after a +CG on 16 July 2004 at 23:07:03.848 UT and the respective MF signal spectrogram. The horizontal lines in the spectrogram are due to medium wave radio transmissions and the vertical line at the triggering time ( $t = 0$ ) is the return stroke MF signal. The lightning characteristics are shown at the right of the VLF plot (distance and azimuth relative to AVF). The 4 panels at the bottom of the figure are the normalized filtered radio signals (normalization with the value of the signal in the milliseconds preceding the return stroke). The radio frequency and the azimuth of the transmitter, with respect to the location of AVF, are given in the respective panels. Strong absorption of the MF radio waves is observed about 1 ms after the lightning electromagnetic signal at time 0 in the top figure. This absorption affects radio propagation at different azimuths and distances from the lightning, indicating a very large disturbed area ( $\sim 200$  km radius) (Farges et al. 2007)

been studied by Farges et al. (2007) based on the effects of about 4000 return strokes on nine chosen radio transmission signals. The main characteristics of the perturbation are a mean peak attenuation of 12 dB, an onset time of less than 1 ms, and duration from 3 to 10 ms. Peak attenuations and durations are roughly proportional to the lightning-induced

peak current. The perturbations have a circular structure centred over the return stroke and reach distances as large as  $\sim 250$  km. Almost all return strokes having a peak current higher than 60 kA perturb radio signals passing through the ionospheric D region at less than 250 km from the lightning.

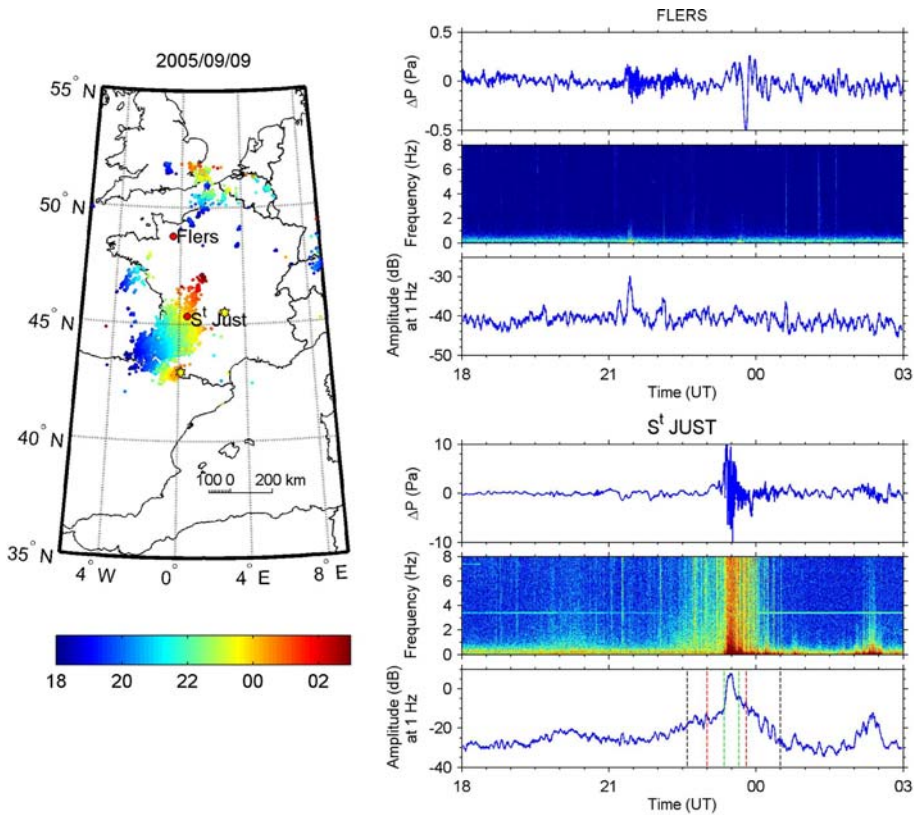
The fading of the MF signals is consistent with impulsive heating of the electrons at  $\sim 90$  km altitude to 4–14 eV by the EMP from the discharges, followed by cooling of the electrons with a time constant (from 0.1 to 100 ms) that matches the recovery time of the signals. The absorption perturbation at MF reflection heights from EMP ionization must be relatively weak compared to that produced by EMP heating (20 dB below) because the recovery time for electron density enhancements is long (from 1 to 100 s, typically 10 s) in comparison with the MF signal recovery time. These perturbations are the MF radio signature of the EMP effect on the lower ionosphere, just as elves are their optical signature. This study shows the capability of this technique to monitor heating changes in the lower ionosphere, even for weaker EMPs that do not lead to enhanced ionization; it complements the VLF technique described above.

## 4 Infrasound from Lightning and Sprites

Infrasound is an acoustic pressure wave with frequencies below the human hearing threshold at  $\sim 20$  Hz. The infrasound stations used during the campaigns are composed of four microbarometers (MB2000) that measure the relative pressure with a sensitivity of  $10^{-3}$  Pa, a band pass of 0.01–27 Hz, and a sampling frequency of 20 Hz. The measurements are used to identify natural infrasound sources and develop discrimination methods as part of the Comprehensive Test Ban Treaty verification procedures. A permanent station is located at Flers in the North West of France and a temporary station was set up in 2005, in the South West of France at St Just, some 400 km South of Flers to deduce the attenuation of infrasound with distance from the lightning source. Several natural and also anthropogenic signals have been measured by infrasound stations, including volcanic eruptions, aurorae, ocean swell, thunderstorms and meteoritic explosions (Blanc 1985), or shock waves from supersonic aircraft (Le Pichon et al. 2002).

### 4.1 Infrasound from Lightning

The thunderstorm of 9 September 2005 provided typical infrasound signals. The storm moved North East from the Bay of Biscay, passing just over the St Just station at around 23:00 UT. The observations are shown in Fig. 11. The right hand part of the figure shows the pressure signals recorded at both stations from 18:00 to 03:00 UT. At St Just, the pressure reaches 20 Pa peak-to-peak when the storm was just over the sensors, whereas it never exceeds 1 Pa peak-to-peak at Flers. The amplitude of the infrasound produced by lightning is higher than the mean signal level when lightning is close to the station. The lightning-induced infrasound amplitude is lower than the noise when it is located 50 km or more from the sensor. When the storm is just over the St Just station, the amplitude at 1 Hz is 40 dB higher than the noise level. The spectrum of the lightning-induced infrasound is very flat from 0.1 to 10 Hz. These measurements show that infrasound produced by lightning and measured at St Just is not detected at Flers, 400 km away. Other signals, recorded during a thunderstorm close to the station and having very large and continuous azimuth dispersion (more than  $100^\circ$ ), are presumably associated with intracloud (IC) lightning over an extended area, termed spider lightning.

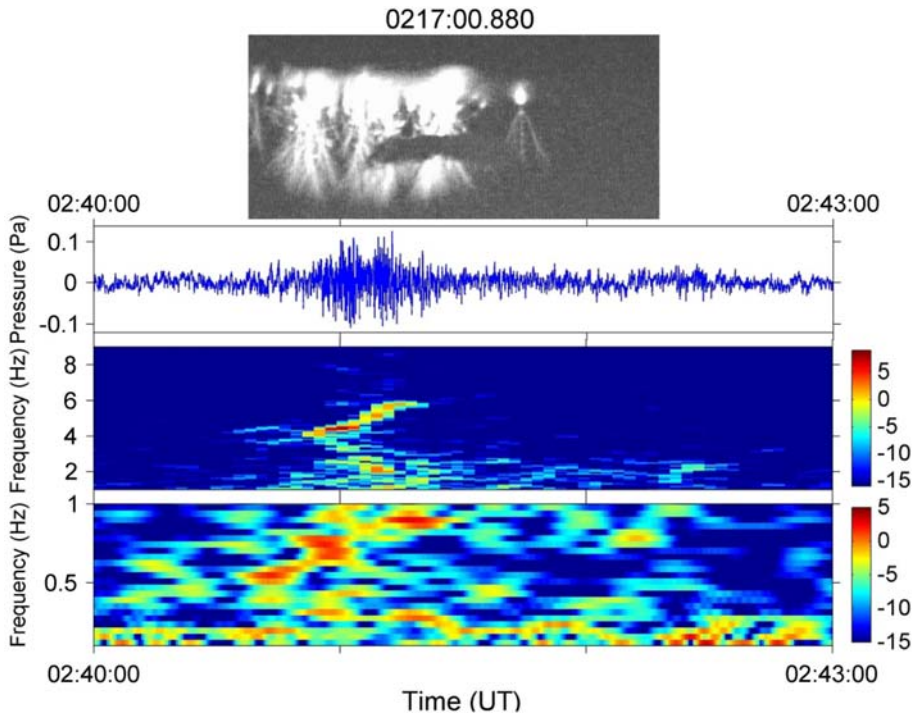


**Fig. 11** (Left) Map of lightning recorded on 9 and 10 September 2005, from 18 to 03 UT, with the locations of the Flers and St Just infrasound stations (red circles) and of camera sites (yellow stars) shown. (Right) Overpressure signal, its spectrogram and the amplitude at 1 Hz recorded at Flers (top) and St Just (bottom). On the St Just amplitude at 1 Hz plot, the vertical dashed lines indicate the mean distance of the storm: less than 50 km (black), 20 km (red) and 10 km (green)

#### 4.2 Infrasound from Sprites

Infrasound waves from sprites were suggested by Liszka (2004), who had noticed a chirp signature in the spectrograms of infrasound detected by a network of sensors in Sweden already in 1995 (Liszka and Hobarra 2006). Unfortunately, no simultaneous sprite images were recorded to confirm or refute this hypothesis. These came for the first time during the *EuroSprite* 2003 campaign where both types of measurements were made (Farges et al. 2005). Examples of the data obtained are shown in Fig. 12. Considering the propagation time of infrasound waves from the source using a ray-tracing model, a clear correlation was found between infrasound signatures and optically identified sprites by cameras at Observatoire Midi-Pyrénées, and over Corsica, to distances up to 1000 km from the Flers infrasound station. A model study shows that a vertically extended cylindrical volume with a radial dimension of the order of several tens of meters placed in the mesosphere and heated by 5°K, consistent with previous analyses (Pasko et al. 1998), is sufficient to explain the amplitude of infrasound from a sprite (Pasko and Snively 2007). The result supports the ideas advanced by Farges et al. (2005) that the chirp-shape can be explained



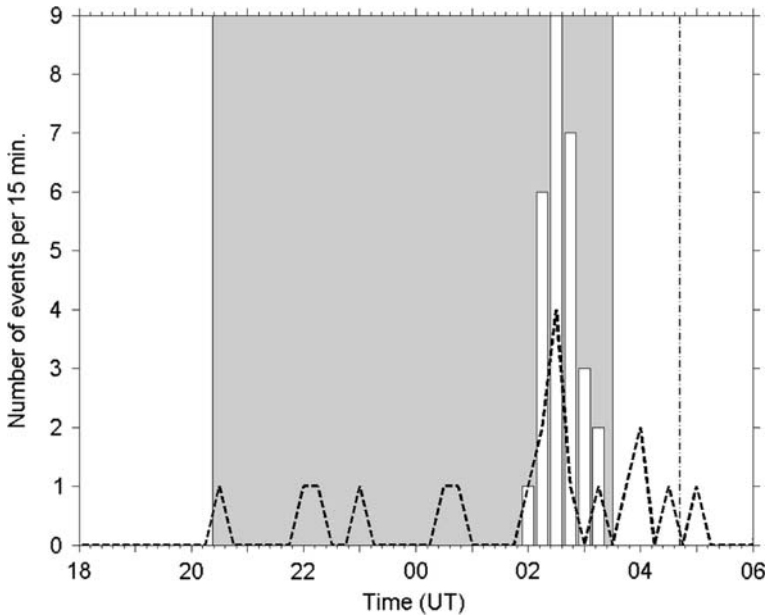


**Fig. 12** Image of sprites observed from Observatoire Midi-Pyrénées at 02.17:00 UT on 21 July 2003, and the characteristics of the associated infrasound signals. The infrasound spectrograms are given for two ranges, 1–9 Hz (upper) and 0.1–1 Hz (lower), with the colour scales being in dB. The infrasound signal was observed at the Flers station at 02.40:50 UT. The propagation delay of 23 min is consistent with a source to receiver distance of about 400 km, computed using a ray tracing model (Farges et al. 2005)

by the horizontal size of the sprite. The sprite-correlated infrasound signatures are characterized by durations of several tens of seconds, which are well correlated with the physical (horizontal) dimensions of sprites observed optically.

During the night of 21 July 2003, infrasound chirp signals were detected after the cameras were shut down because of the approaching dawn (Farges et al. 2005). Several chirp-shaped infrasound waves were measured from 3:50 to 6:00 UT. For the cameras, dawn starts when the Sun is  $12^\circ$  below the horizon, i.e. at 3:50 UT on this date. However, D-region ionization starts to increase from 04.45 UT. The conductivity of the middle atmosphere is thus expected to remain the same up to this time. The infrasound station observed 3 sprites up to about one hour after illumination of the D-region (from 4:45 to 6:00 UT). After 6:00 UT, the conductivity had reached such a level that the electric field produced by lightning is below the threshold for breakdown. The result shown in Fig. 13 is consistent with the conclusions of Stanley et al. (2000) based on ELF signatures of sprites near sunrise.

It is concluded that the detection of sprites by infrasound is an attractive alternative to optical detection, as it is not limited by clear viewing conditions or by the absence of daylight. The infrasound signature is so characteristic that it makes automated detection of sprites possible (Ignaccolo et al. 2008) (see Sect. 6). The automatic detection of sprites in infrasound data opens up the possibility of an interesting scientific use of data obtained by the infrasound stations of the Comprehensive Test Ban Treaty monitoring network.



**Fig. 13** The number of sprites per quarter of an hour observed from Observatoire Midi-Pyrénées on 20 and 21 July 2003 (full line, vertical bars) and the number of infrasound chirps per quarter of an hour measured at Flers (dashed line). The grey zone indicates night-time, using nautical sunset/sunrise times; the vertical dash-dot line marks the time of sunrise in the D-region

## 5 Electromagnetic Radiation from Lightning and Sprites

The electromagnetic (EM) radiation from electric discharges in the atmosphere can be observed to large distances of  $\sim 1000$  km or more. The frequency bands of the radiation emitted are determined by the time scales of different processes acting in the discharges, making EM radiation a powerful tool for studying lightning and TLEs. Since a sprite is triggered by a lightning discharge and therefore occurs almost simultaneously with it, it is a challenge to separate the radio signature of a sprite from that of the causative lightning discharge.

### 5.1 Radiation at ULF

The ultra low frequency (ULF) ( $\sim 0.1$ – $10$  Hz) signals of +CGs that generate sprites and those that do not were investigated in a one-to-one case-study and in a statistical approach which included 240 strong +CG events. The ULF observations were at the southernmost Finnish station of Nurmijärvi (NUR), whose average great circle distance from the European thunderstorms studied during the *EuroSprite* campaigns is  $\sim 2000$  km.

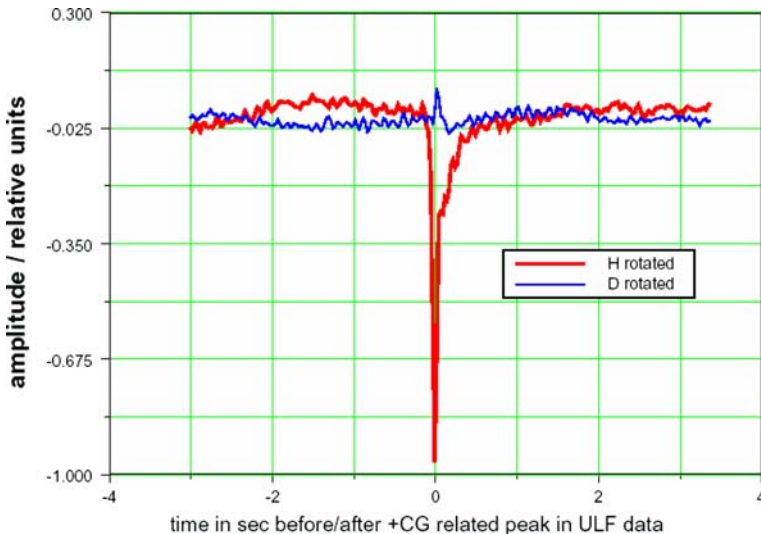
The pulsation magnetometer used here is optimized for the observation of long period geomagnetic pulsations (period range  $\sim 200$  to  $0.1$  s); it is almost insensitive to broad-band electromagnetic emissions which originate from processes characterized by time scales from microseconds to some tens of ms, the dominant time scales of lightning discharge and sprite processes (Füllekrug et al. 2006a). In spite of this fact a large proportion of lightning discharges produce a detectable trace in the recordings of pulsation magnetometers around the world (Fukunishi et al. 1997). This is because, in association with the lightning flash,

there are secondary and/or tertiary processes having much longer time scales which emit electromagnetic waves as well. This is a fundamental point when searching for a unique signature in the ULF-range to sprite-associated lightning flashes.

From the original Météorage data set of cloud-to-ground lightning discharges occurring over France, a subset was created (list 1) of only those +CG and –CG events which stem from the same thunderstorm centres as observed optically by the CAL cameras. The CAL list of sprite-associated +CG events is a subset of list 1 (termed list 2). Using both lists, those +CG events which are followed by a sprite might be distinguished from those which are not. The general properties of the ULF response to sprite-associated lightning flashes were then collected and analysed, and their properties compared with those of a control group not associated with a sprite.

The comparison did not reveal any unique and identifiable ULF signatures relating to sprites (Bösinger et al. 2006). The analysis resulted, however, in categorizing a collection of ULF responses to +CG (and actually also to –CG) discharges, which has not—to our knowledge—been compiled so far in a similar systematic fashion. Some of the properties of ULF responses to +CGs have already provoked subsequent investigations (Shalimov and Bösinger 2006, 2008). One typical result of the superposed epoch analysis is shown in Fig. 14. All responses represent single, isolated, strong events, the so called Q-bursts (Ogawa et al. 1967), which exceed the average natural background noise level (primarily caused by world-wide thunderstorm activity) by at least an order of magnitude.

It is essential to be certain of the selection of non-sprite-associated +CGs forming the control group; this is not trivial since sprites can be missed for various reasons such as



**Fig. 14** The result of superposing the ULF response at NUR of 32 clean and isolated +CG events which were all associated with sprites; the plot for the control group looks similar. The data have been corrected for the frequency and phase response of the instruments, and the coordinate system was rotated such that the response maximizes (in absolute terms) in one component (H-component). This was achieved by rotating the horizontal H-D plane of the Cartesian coordinate system (H, D, Z) around its vertical Z-axis. The rotation angle is chosen such that the disturbance minimizes for the D-component (East–West orientation). The zero time represents the data point with the minimum (or peak) value in the rotated H-component (North–South orientation). The individual time series in the superposition were normalized to this minimum. The small deflection in this component indicates an almost totally linearly polarized disturbance vector. The amplitudes of these Q-bursts were of the order of 10 pT (Bösinger et al. 2006)

clouds, light pollution, low emission intensities or simply because the camera's field of view (FOV) does not cover the entire thunderstorm cell with active lightning. This difficulty was overcome by using the same data set as Haldoupis et al. (2004) and Mika et al. (2005); the one-to-one association found by these authors already implied the pinpointing of every +CG within the camera's FOV. Moreover, their various VLF links crossing the thunderstorm regions studied provided additional information on sprite occurrence.

## 5.2 Radiation at ELF

Extremely Low Frequency (ELF) electromagnetic radiation, ranging from 8 Hz to 3 kHz, from particularly intense lightning discharges and sprites can be used for remote sensing of both types of discharge. Since the ELF radiation propagates over long distances within the Earth-ionosphere cavity, the recordings can be used to determine global occurrence rate of sprites (see Sect. 6).

Sprites are produced by +CG discharges (Boccippio et al. 1995), with the lightning having a strong continuing current (Reising et al. 1996), and they cause Earth-ionosphere cavity (or Schumann) resonances (Füllekrug and Reising 1998). The sprite appears  $\sim 5$  ms after the causative lightning discharge and radiates ELF electromagnetic waves (Cummer et al. 1998) as a result of mesospheric breakdown (Füllekrug et al. 2001).

To study the influence of the natural atmospheric variability on the generation of sprites, a model of lightning electric fields above thunderstorms was developed (Füllekrug 2006). The model uses typical parameters of particularly intense positive lightning discharges which cause sprites. For this, the charge transfer within intense lightning discharges needs to be determined.

A mean lightning current waveform was inferred from the recordings of 52,510 intense lightning discharges (Füllekrug et al. 2006b) made with a global network of ELF magnetometers (Füllekrug and Constable 2000). It was found that the lightning current initially decays with a time constant of  $\sim 2$  ms and that it lowers  $\sim 60$  C from cloud to ground within the first  $\sim 10$  ms of the discharge. The subsequent continuing current exhibits a decay time constant of  $\sim 40$  ms and lowers  $\sim 170$  C from cloud to ground within the next  $\sim 100$  ms. The current associated with the total charge transfer of  $\sim 230$  C from cloud to ground deposits electrical energy in the middle atmosphere via quasi-static (Joule) heating. The energy deposition is dominated by the lightning continuing current and is  $\sim 10^{-5}$  Jm $^{-3}$  at 30 km altitude. The associated peak electric field in the mesosphere is generally not sufficient to produce conventional breakdown, such that the natural variability of other physical parameters seems to be important for sprite production. The least known parameters in the model are the altitude and the horizontal extent of the lightning discharge which causes a sprite, and also the atmospheric conductivity which can vary by an order of magnitude, e.g., between day and night. The atmospheric conductivity above thunderstorms may be strongly influenced by the electrical activity within the thunderstorms themselves through quasi-static heating.

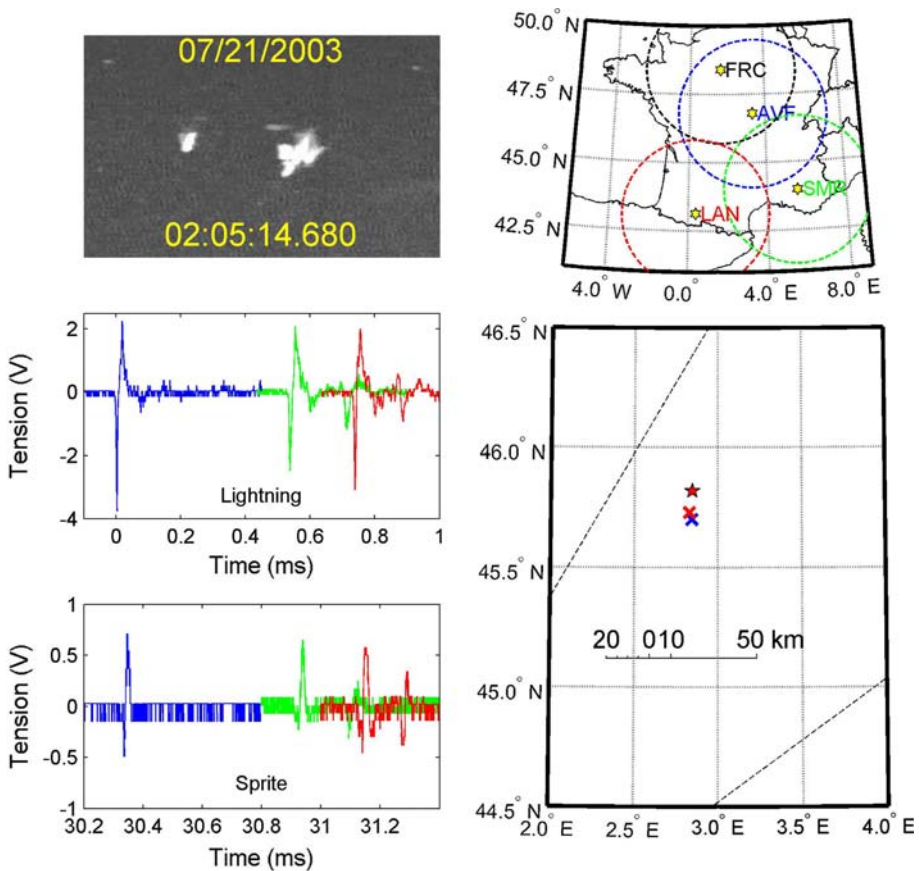
## 5.3 Radiation at Higher Frequencies

In 2003 and 2004, four stations were deployed in France to observe electromagnetic radiation from lightning and sprites. The objective was to measure the vertical electric field across the frequency ranges from VLF (3–30 kHz) to LF (30–300 kHz), MF (0.3–3 MHz) and HF (3–30 MHz), and to detect the possible burst of radiation associated with the runaway breakdown process (Roussel Dupré et al. 2002; Gurevich and Zybin 2004). In 2003, possible

radio emissions were correlated with observations of CG discharges made by Météorage with optical observations of sprites made from the Observatoire Midi-Pyrénées.

Weak HF emissions, i.e. enhancements of the signal from 0.5 to 3 MHz, were measured during the time interval when sprites were observed. The system recorded HF signals for 58 of 133 sprites observed in 2003 (Appendix 1). To prevent any confusion between radiation from lightning and from sprites, we select only events for which no CG discharge is observed during the time interval where a sprite is observed with the camera (20 ms exposure time). 8 events were kept after this selection, with HF bursts being associated with sprites in 5 cases, or  $\sim 60\%$ .

In several cases, HF emissions were recorded simultaneously with VLF emissions which are presumably either from the sprite or from intracloud (IC) lightning, which has been associated with carrot sprites (van der Velde et al. 2006). With at least three receiver stations, it is possible to localize the source of the VLF signals. Figure 15 shows the image of a sprite



**Fig. 15** VLF radiation from lightning and/or sprite. (Left, top) sprite image at 02:05:14.668 UT 21 July 2003, (left, middle) waveforms of the lightning, observed at the station AVF in blue, at SMR in green and at LAN, (left, bottom) waveforms of the sprite with the same colour presentation, (right, top) location of Commissariat à l’Energie Atomique (CEA) stations, (right, bottom) lightning localisation by Météorage (blue cross), and lightning (red cross) and sprite (red star) localisations using CEA station data. The dashed lines represent the camera field of view (FOV)

taken on 21 July 2003 at 02:05:14.680 UT, from 12 to 32 ms after the parent lightning. Also shown is the VLF signal induced by the parent lightning and a second signal 30 ms after that, within the time interval of the optical observation. The triangulation calculation locates the lightning (red cross) as being only 1 km away from the location given by Météorage (blue cross), which corresponds to the precision of the Météorage localisation. The location of the second VLF signal is 10 km away (red star), a separation distance which is often observed. The second VLF signal is due either to a CG with a very weak peak current or to an IC discharge, both undetected by Météorage, or to the sprite itself.

## 6 The Global Rate of Sprites

An assessment of the planetary rate of sprite events is necessary to determine the global impact of sprites on the properties of the atmosphere such as perturbations to the chemical composition, mesospheric conductivity and the global electrical circuit. However, almost 20 years after the first sprite observations, the rate of sprite occurrence over the Earth is yet to be reliably determined. Several estimates of the planetary rate of sprite events have been made. These estimates range from 0.5 to 33 sprites per minute (Sato and Fukunishi 2003; Yair et al. 2004; Blanc et al. 2004).

### 6.1 Algorithm to Estimate the Global Sprite Rate

A new formula to find the planetary rate of sprite events has been developed based on remote sensing strategies, for various detectors of optical, infrasound and radio wave sprite signatures (Ignaccolo et al. 2006). The formula uses the number of detected sprites, the detection efficiency and the false alarm rate of each detector along with spatial and temporal effectiveness functions. The formula estimated an average planetary rate of sprite events of  $\sim 2.8$  per min, with an accuracy of a factor  $\sim 2$ – $3$ . The proposed formula is general, and can be used to calculate the occurrence rate of any physical event detected by remote sensing.

An important element for determining the planetary rate of sprite events is the newly discovered infrasound signature of sprites for which an automated detection algorithm has been developed (Ignaccolo et al. 2008). Three different methods were compared to detect the sprite infrasound signature (chirps) automatically in random background noise: (1) the standard deviation detector uses the signal-to-noise ratio of the measurements, (2) the spectral covariance detector uses the spectrogram to detect the spectral features of chirps in background noise, and (3) the “Complex Analysis of Sequences via Scaling AND Randomness Assessment” (CASSANDRA) detector (Allegrini et al. 2002; Allegrini et al. 2003) uses diffusion entropy analysis to detect periodic patterns such as chirps in noise. All three detectors were applied to the observed infrasound recordings to detect the chirps produced by sprites. It was found that the CASSANDRA detector provided the best trade-off between the false alarm rate and the detection efficiency for automated sprite detection. With optimised parameters, the algorithm achieved a null false alarm rate with a detection efficiency of  $\sim 66\%$ .

## 7 Perturbations to the Global Atmospheric Electric Circuit

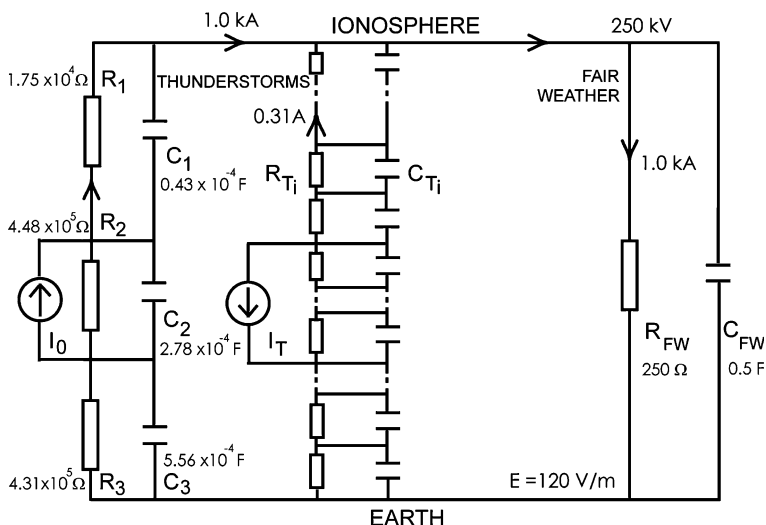
Thunderstorms around the world are one important generator which powers the Earth’s electric circuit. Thunderstorms, generally being primarily positive at the tops of clouds,

drive an electric current up to the electrically good conducting ionosphere, which distribute the currents and the associated electrical potential globally. Currents flow back to the Earth in the fair weather regions, where these currents are of the order of  $10^{-12}$  A/m<sup>-2</sup> and the electric field at ground level is  $\sim 120$  V/m, directed downwards. The Earth's global atmospheric electric circuit is reviewed within the wider field of planetary atmospheric electricity in a comprehensive new book (Leblanc et al. 2008). In this paper, studies of the global electric circuit have been undertaken to estimate the effects of lightning discharges and sprites, as outlined below.

### 7.1 The Circuit Model of a Thunderstorm and the Atmospheric Electric Circuit

Using the commercially available electrical engineering software package PSpice (Personal SPICE—Simulation Program with Integrated Circuit Elements), a novel model for the operation of the global atmospheric electric circuit has been developed. Shown in Fig. 16, this consists of a current generator  $I_0$  representing the  $\sim 1000$  thunderstorms that are active at any one time. A “battery” between the thunderclouds' negative charge at 5 km altitude and the positive charge at 15 km drives an upward current which maintains the potential of the ionosphere at 80 km altitude at up to  $\sim 300$  kV with respect to the Earth. The values of the resistors  $R_1$ ,  $R_2$  and  $R_3$  are derived from a model of the height profile of the electrical conductivity of the atmosphere such that, using Ohm's law, the current flowing in the circuit is  $\sim 1000$  A.

One particular thunderstorm is represented (see the central part of Fig. 16) by a current generator  $I_T$ , it being assumed that there are 80  $R$  components and 80  $C$  components, each representing a 1 km thick layer of the atmosphere (so that  $i$  goes from 1 to 80). The values of the capacitors  $C$  are derived from standard electrostatic theory for plane capacitors. Both the background, with a full literature survey, to this study and complete details of the



**Fig. 16** Model circuit diagram (similar to that presented by Rycroft et al. 2007), representing 1000 thunderstorms worldwide generating  $\sim 1$  kA flowing in the global atmospheric electric circuit through the fair weather resistance  $\sim 250$  ohms. The model of an isolated thunderstorm generating +CG discharges is shown in the centre

model, together with its justification, have recently been presented by Rycroft et al. (2007). This model has, in fact, been extended to include currents generated by electrified shower clouds which are believed to be as important as those generated by thunderstorms.

The good conducting ionosphere is assumed to be an equipotential surface, as is the good conducting Earth. Between the two, in regions remote from thunderstorms and electrified shower clouds, and termed the fair weather region (almost 98% of the Earth's surface), there is a downward current, with a current density of  $\sim 2 \text{ pA/m}^2$ . Near the Earth's surface there is a downward electric field, or a positive potential gradient (as height increases), of  $\sim 120 \text{ V/m}$ .

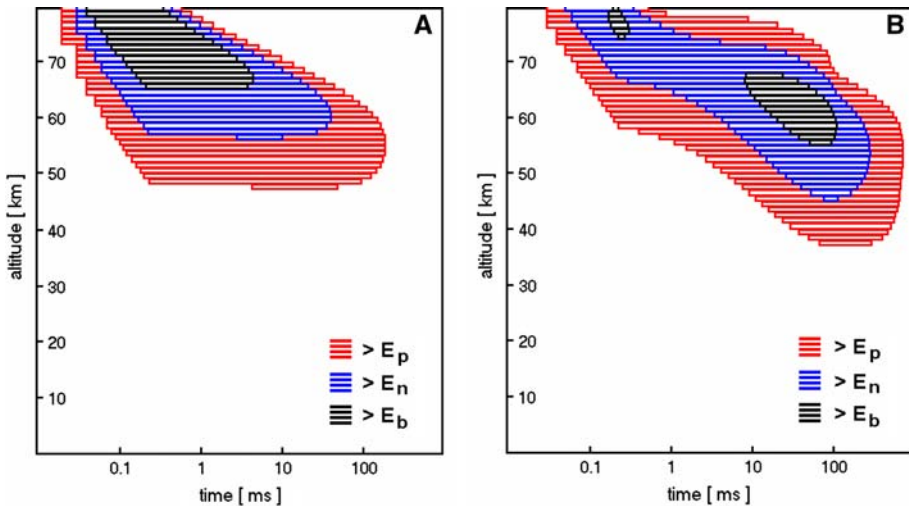
To simulate lightning discharges occurring in the atmosphere, voltage-controlled switches are included in the circuit in parallel with the thunderstorm  $T$ . A negative cloud-to-ground lightning discharge having the median characteristics of observed lightning has been satisfactorily simulated. By inverting the polarity of the thunderstorm, a positive cloud-to-ground (+CG) stroke has similarly been simulated.

## 7.2 Some Results

Two different +CG discharges from the region at 140 MV at 5 km altitude have been modelled, to examine the differences in the effects on the mesosphere which may result in different kinds of upper atmosphere discharges such as sprites. Observations suggest, as mentioned in Sect. 2.2, that "column" sprites occur several ms after a +CG discharge with a large ( $\sim 100 \text{ kA}$ ) return stroke current, while "carrot" sprites might be delayed by some tens of ms after a smaller ( $\sim 30 \text{ kA}$ ) return stroke current with a strong continuing current caused by the redistribution of charges by preceding inter-cloud (IC) lightning activity. Here, the first modelled +CG discharge (A) as a return stroke with 110 kA peak current lasting 0.2 ms, while the second (B) is a +CG 30 kA return stroke lasting 0.2 ms and a continuing current of 2 kA for 90 ms. During the first discharge 20 C is transferred to ground, whereas the second transfers 70 C in total. With their image in the conducting Earth, these correspond to a dipole moment, termed the charge moment, of 200 and 700 C km, respectively. The +CG discharges reduce the ionospheric potential everywhere by  $\sim 10 \text{ V}$  and  $40 \text{ V}$ , respectively, in  $\sim 20 \text{ s}$ . The ionospheric potential recovers with the  $\sim 250 \text{ s}$  RC time constant of the fair weather circuit. If there were to be ten such discharges around the world in  $\sim 20 \text{ s}$ , as might well occur, the ionospheric potential would be reduced by between 100 and 400 V.

Analysis of the changes of the vertical electric field above the thundercloud caused by the modelled +CG discharges reveals differences which can explain the occurrence of different types of sprites. Figure 17 shows the extent in altitude and time of regions above the thundercloud where the vertical electric field exceeds a certain threshold field, namely the breakdown threshold  $E_b$  which has to be exceeded for the creation of streamers (see Sect. 8), and the thresholds required for sustaining the propagation of positive and negative streamers,  $E_p$  and  $E_n$ , respectively (Pasko 2006). For the return stroke with the larger current, the region where the breakdown threshold field was exceeded occurs at altitudes from  $\sim 80$  to 65 km within 0.01 to 1 ms after the current starts. A similar effect, but of much shorter duration (only 0.04 ms) and lesser altitude extent ( $\sim 80$  to 75 km) appears after the 30 kA return stroke, but breakdown occurs again later at lower altitudes of  $\sim 75$  to 65 km. These modelled changes of electric field agree with the scenario that columniform sprites are created at higher altitudes (80 to 75 km), with positive streamers propagating downwards, and with carrot sprites being more delayed and originating at lower altitudes ( $\sim 65 \text{ km}$ ), with both positive and negative streamers propagating



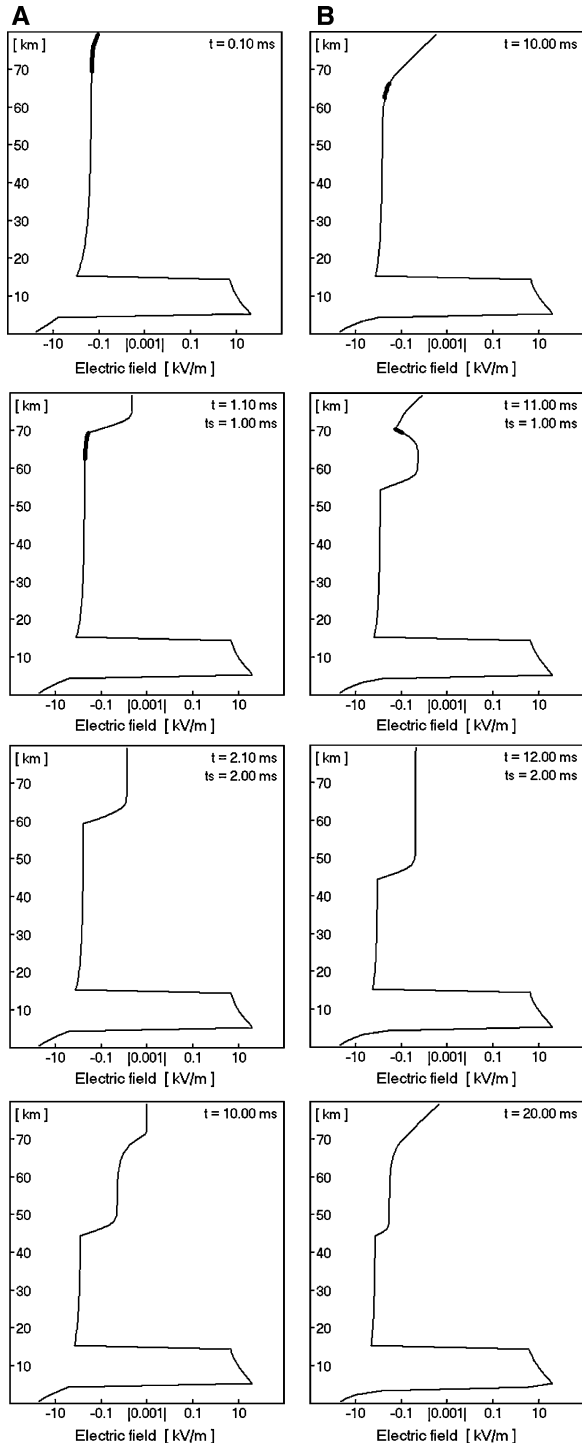


**Fig. 17** Diagram showing the altitudes and times where and when three different breakdown threshold fields ( $E_b$ ,  $E_n$  and  $E_p$ ) are exceeded following +CG discharges A (left) and B, with a continuing current (right). The electric field in both cases is always directed downwards. In the region where the threshold field  $E_b$  is exceeded, streamers can be created and propagate. Positive streamers propagate downwards where the magnitude of the electric field is greater than  $E_p$ ; negative streamers propagate upwards where the magnitude of the electric field is greater than  $E_n$ . During discharge A,  $E_b$  is exceeded  $\sim 0.1$  ms after the onset of the return stroke at the ionosphere, and down to  $\sim 65$  km after a few ms. During B, between 0.17 and 0.21 ms after the return stroke onset, this region is created at  $\sim 78$  km where a sprite halo may be produced. Then similar conditions are created  $\sim 10$  ms after the return stroke onset at  $\sim 65$  km; a situation which lasts for several tens of ms at altitudes  $>55$  km. Also at this time the threshold for the formation of positive and negative streamers is exceeded over a wide range of altitudes, so that the sprite can be produced by both downward and upward propagating streamers over a height range of  $\sim 30$  km

downwards and upwards, respectively (Stanley et al. 1999; McHarg et al. 2007, and references therein).

The sprite is modelled as a region of significantly increased conductivity,  $\sim 10^{-7}$  S/m, expanding downwards to 50 km altitude with a velocity of 10 km/ms, again using switches. Sprite A (column-like) starts with the closing of the first switch at 80 km, each switch representing the volume above the thunderstorm of thickness 1 km, 0.1 ms after the 110 kA return stroke. For the carrot-like sprite B, the switch at 65 km closes first, 10 ms after the 30 kA return stroke and during the 2 kA continuing current. In this case the upward propagation of a sprite was modelled by subsequently closing switches from 65 to 80 km, starting at 10.5 ms, thereby contributing to the spatial and temporal development of the sprite. One ms after sprite A starts it has developed down to 70 km, and a further 1 ms later (at sprite time,  $t_s = 2$  ms) down to 60 km. On the other hand, sprite B developed down to 65 km in 1 ms, and a further 1 ms later (at sprite time  $t_s = 2$  ms) down even to 45 km and up to 80 km. These features are clearly seen in the middle four panels of Fig. 18. Breakdown occurs where and when the profile is shown in bold; large electric fields remain in this region for some tens of ms. Both sprites A and B reduce the ionospheric potential globally by only  $\sim <2$  V. If there were to be two sprites in 20 s, the potential of the ionosphere would be expected to change by  $\sim <3$  V. Associated with the sharp vertical changes of electric field there are charge layers, as shown in the various panels of Fig. 18. These are especially marked at the top and bottom of the thundercloud and also at the top and bottom of the sprite.

**Fig. 18** Changes of the vertical electric field above a thundercloud caused by a model +CG discharge and a subsequent sprite are shown: the left hand panels are for discharge A and sprite A, whereas the right hand panels are for discharge B (with a continuing current) and sprite B, at various times after the return stroke starts. The bold line shows where  $E_b$  is exceeded. Sprite A, initiated at 80 km at 0.1 ms after the return stroke field reaches 80 km, propagates downwards, reaching down to 45 km in 3.5 ms. Sprite B is initiated at 65 km, after 10 ms, and propagates both downwards to 45 km and upwards to 80 km, in 2 ms. The occurrence of a sprite reduces the electric field above the thunderstorm at sprite altitudes. The situation depicted here is more complicated than that shown by Rycroft et al. (2007)



During the *EuroSprite* campaigns, events similar to sprite A were observed on 26 September 2005 at 01:49:15.291–332 UT (on one video frame having a duration of 41 ms) and at 03:10:29.186–227 UT, and also on 17 August 2006 at 22:58:46.203–244 UT. Sprites similar to sprite B were recorded on 29 July 2005 at 01:32:51.249–333 UT, on 17 August 2006 at 23:38:07.736–789 UT and on 13 October 2006 at 22:17:00.597–634 UT.

These results indicate that, whilst there can be strong effects above a thundercloud following a +CG lightning discharge, these are not expected to have any effects which might be noted on the global scale. For example, the fair weather electric field would change by only  $\sim 0.01\%$  following one of these model +CG discharges, or by less than 1 part in  $10^5$  after two model sprites occurring in 20 s.

## 8 Particle Kinetic Simulations of the Sprite Discharge

Electric discharges in gases can take a number of forms. The most well known discharge in the atmosphere is lightning between electrically charged regions of a cloud or from a cloud to the ground. Although lightning has been studied for many years, there are still new aspects of this natural phenomenon. While observations of X- and  $\gamma$ -rays associated with the charging of clouds dates back to the 1980s (Parks et al. 1981; McCarthy and Parks 1985; Eack et al. 1996), the discovery of X- and  $\gamma$ -rays in association with lightning is only a few years old (Moore et al. 2001). It is thought that X- and  $\gamma$ -rays are bremsstrahlung radiation from energetic electrons accelerated in the electric fields of clouds and lightning (Gurevich et al. 1992, 1999; Roussel-Dupré et al. 1994). The study of this phenomenon has gained momentum with the discovery from space of bursts of X- and  $\gamma$ -rays observed by satellite from regions of the atmosphere above thunderstorms (Fishman et al. 1994; Smith et al. 2005).

Laboratory studies of electric discharges date back to the late nineteenth century, and many of the fundamentals of the discharge process are now well established (Raizer 1997). The characteristics of laboratory discharges depend on the medium in which they occur and the boundary conditions such as electrode geometry, material, applied voltage, etc. At ground pressure in air, the electrical breakdown voltage is  $\sim 3$  MV/m. The breakdown voltage is inversely proportional to the gas density  $N$  in which the breakdown occurs, and often the “reduced” electric field  $E_r = E/N$  is used to characterise the electric field in a gas (The unit of this reduced electric field is the Townsend (Td), where  $1 \text{ Td} = 10^{-21} \text{ Vm}^2$ ).

Discharges may develop from a single electron which spawns additional electrons by collisionally ionizing the gas, creating an ionization avalanche. When the avalanche has grown in magnitude to the point where the space charge fields are significant relative to the background electric field, the avalanche can develop into a streamer. The streamer creates an ionization path through the gas that may allow additional charge to flow into the streamer path from the electrodes. The streamer then becomes a leader, with a high electron density and a small electric field within the highly conducting leader region. The development of this chain process may occur in steps; a streamer propagates a limited distance and stops until the streamer region is again filled with electrons, turning the region into a leader and bringing with it the electric potential, and then continues from the tip of this region to send forth a new streamer. This process is well known in lightning, and is termed stepped leader propagation (Rakov and Uman 2003).

A streamer may propagate from the anode towards the cathode, i.e. in the direction of the electric field, and one speaks of a positive streamer, or against the electric field, which

is a negative streamer. It is generally thought that the downward propagating elements of sprites to low altitudes are positive streamers while upward propagating elements are negative streamers. The upward and downward propagation of streamers are seen in high-speed video recordings of sprites (Cummer et al. 2006a; McHarg et al. 2007). It is interesting to note the fundamental difference between negative and positive streamers. For negative streamers electrons are accelerated in the direction of the streamer propagation and thus create the additional avalanche ionization needed ahead of the streamer. For positive streamers, electrons are accelerated against the propagation direction and a second ionization process is needed. This is either ionization from background electrons if the gas is partially ionized or ionization from photons emitted from the high electric field region of the streamer head. In the atmosphere, the energy of photons emitted by the nitrogen molecule—when excited—is sufficient to ionize the oxygen molecule (Zheleznyak et al. 1982). This reaction path is a fundamental process of atmospheric electric discharges.

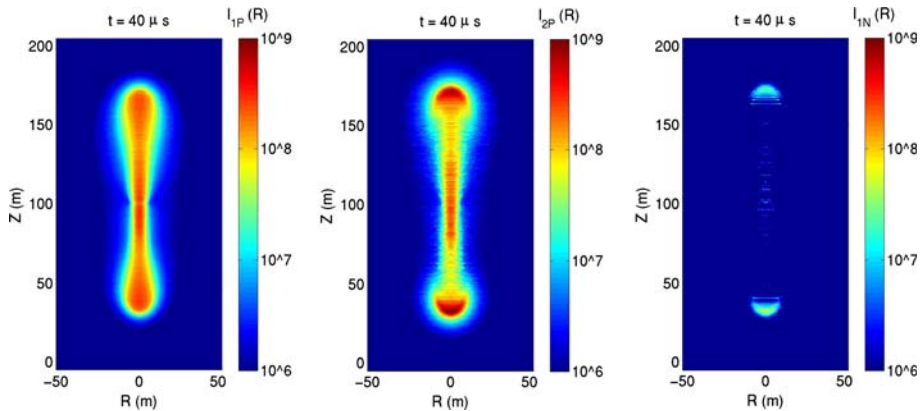
Computer simulations of discharges are difficult because very high spatial resolution is required in the regions around the streamer tips where the space charge fields vary strongly (Ebert et al. 2006). Likewise, it is a problem to resolve the electron kinetics because of the exponential growth of the number of electrons in a discharge. For these reasons, only limited aspects of the problem have been simulated. Often, simulations are carried out in one dimension (1D) and/or in a fluid description. However, advances in computer speed, memory size and numerical techniques have allowed increasingly sophisticated simulations (Arrayás et al. 2002; Liu and Pasko 2004; Montijn et al. 2006). Some recent particle simulations of streamer development at sprite altitudes are presented here.

## 8.1 The Simulation Model

A computer particle code has been developed to study particle kinetics of electric discharges in air. The code is in 2D axi-symmetrical coordinates, simulating the development of an electron avalanche, its transition into a streamer and the propagation of the streamer. The code includes photoionization processes that allow the development and propagation also of positive streamers. However, effects of pre-ionisation of the atmosphere or negative ions are not considered, although they may affect positive streamer propagation (Pancheshnyi 2005). The code uses a Particle-In-Cell (PIC) technique with Monte Carlo collisions, where the charged particles are followed in a Cartesian mesh, the electric field is updated at grid-points from the charged particle densities and Poisson's equation, and collisions with neutral air molecules are random according to cross section probabilities (Birdsall 1991). The code is used to study streamer generation and propagation under atmospheric conditions of sprite ignition in the mesosphere and the associated particle acceleration and optical emissions. A detailed description of the code is given in Chanrion and Neubert (2008).

## 8.2 Optical Emissions from Streamers at Sprite Altitudes

By colliding with air constituents, the electrons of a streamer discharge excite the air molecules. Returning to their ground state, the excited air molecules emit photons producing optical emissions from the streamer. The main optical emissions of sprite streamers in the visible range come mostly from the 1st positive, 2nd positive and 1st negative nitrogen bands. From the Monte Carlo part of the code, one can calculate the production rate of excited neutrals along streamer simulation from which the optical emission rates can be derived by solving the transition equations.



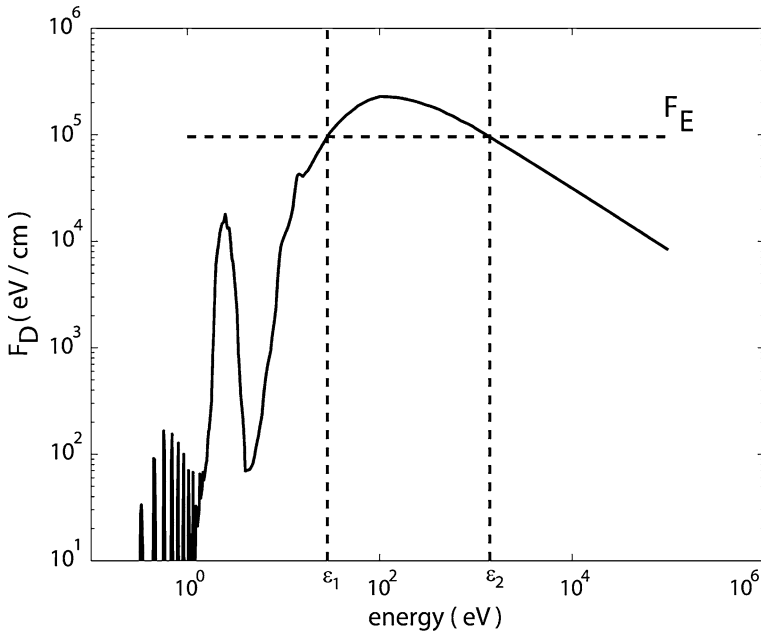
**Fig. 19** Optical emission intensities for a streamer simulation at sprite altitude with the background field  $E_o = 1.5 E_k$ . From left to right panel: nitrogen 1st positive (1P), 2nd positive (2P) and 1st negative (1 N)

Results from a streamer simulation are presented in Fig. 19. Following Liu and Pasko (2004), the streamer is simulated in the mesosphere, at 70 km altitude in a background electric field of 1.5 times the electric field threshold for conventional breakdown and is initiated from a plasma cloud with peak density of  $5 \times 10^5 \text{ cm}^{-3}$  and a Gaussian spatial distribution with a scale of 3 m. The intensities of optical emissions in different spectral bands are shown at the time  $t = 40 \mu\text{s}$ . The results are in good agreement with the fluid simulations of Liu and Pasko (2004). Estimates of streamer brightness from recent high speed imaging of sprite streamers give an average streamer head intensity ranging in the interval  $9 \times 10^8$ – $5 \times 10^{11}$  R (Stenbeak-Nielsen et al. 2007). Our simulation gives an average intensity of  $1.5 \times 10^9$  R, which is within the range of intensities derived from the observations.

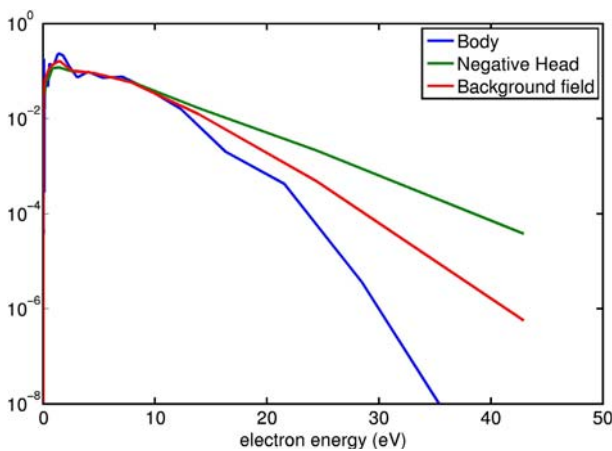
### 8.3 High Energy Electrons from Streamers

By using a non isotropic scattering model for electron collisions with air molecules and a set of cross-sections that goes up to 10 keV electron energy, the code can study the production of high energy electrons by sprite streamers or by streamer regions of lightning leaders as described by Moss et al. 2006. The most interesting point is to understand if the streamer's electric field can accelerate electrons to their runaway regime from which electrons can gain sufficient energy to produce X- and  $\gamma$ -ray emission by bremsstrahlung radiation. The runaway process can be understood by comparing the friction force  $F_D$ , an average force that represents the stopping power due to collisions, and the electric force  $F_E$ , from the electric field accelerating electrons to higher energy. The two forces are represented in Fig. 20 in air at sea level pressure as a function of the electron energy. The two forces are of equal magnitude for two energies  $\varepsilon_1$  and  $\varepsilon_2$  that represent, respectively, a stable and an unstable equilibrium. The runaway regime is defined by the energy  $\varepsilon_2$  above which electrons keep gaining energy from the electric field.

The electron distribution functions corresponding with the simulation presented in Sect. 8.2 at the time  $40 \mu\text{s}$  are shown in Fig. 21. The distribution function for electrons in the body of the streamer are shown in blue (the background electric is partially screened), in the negative streamer head in green (the background field is enhanced) and in the region



**Fig. 20** The frictional force as a function of energy of free electrons in air composed of 80% nitrogen and 20% oxygen at atmospheric pressure at sea level. The horizontal dashed line shows the force on an electron exerted by a background electric field 50% greater than the breakdown threshold electric field. The intersections with the curve mark the electron energies at which the two forces balance. Point 1 is a stable equilibrium and point 2 unstable, allowing electrons to be accelerated to high energies (see text). At energies above 1 MeV the force begins to increase again due to the increasing relativistic mass of the electron (Chanrion and Neubert 2008)



**Fig. 21** The electron distribution function for the streamer simulation presented in Fig. 19. The distribution function is shown for electrons in the background gas dominated by the background electric field, in the streamer body, which is the region with the lowest electric field and in the streamer head where the field is a maximum

surrounding the streamer in red (background electric field). The distribution functions show the effect of the streamer formation on the energy levels reachable by electrons. In the body of the streamer, the electric field is reduced and thus electrons are less energetic than if they were drifting in the background field. In the streamer tip the field is enhanced and the electrons are more energetic. Depending on the magnitude of the background electric field and on the altitude of the streamer, the enhancement of the field in the streamer tip can be high enough to accelerate electrons into the runaway regime. These runaway electrons created by streamers could generate a source of energetic electrons able to produce X- and  $\gamma$ -ray emission by bremsstrahlung radiation.

## 9 Simulations of Perturbations to $\text{NO}_x$ and Ozone by Sprites

Odd nitrogen (N, NO and  $\text{NO}_2$ , together known as  $\text{NO}_x$ ) and ozone are among the most important chemically active species in the middle and upper atmosphere. Ozone controls the dynamic balance of the atmosphere due to its absorption of solar radiation, and  $\text{NO}_x$  affects the concentration of ozone. In the troposphere  $\text{NO}_x$  creates ozone whereas in the stratosphere and mesosphere it destroys ozone. The concentration of  $\text{NO}_x$  in the mesosphere is known to be enhanced by transient events such as auroras and solar proton events (SPEs) (Crutzen and Solomon 1980). Comparisons of  $\text{NO}_x$  concentrations predicted by local models with satellite measurements show that the theoretical understanding of thermospheric  $\text{NO}_x$  concentrations is not complete and that vertical transport by neutral winds must be important (Barth and Bailey 2004; Saetre et al. 2007). Mesospheric  $\text{NO}_x$  is long-lived in the winter polar vortex and can subside into the stratosphere (Seppälä et al. 2007). The main stratospheric source of  $\text{NO}_x$ , however, is oxidation of  $\text{N}_2\text{O}$  from the troposphere (Vitt and Jackman 1996).

Thunderstorms can contribute to  $\text{NO}_x$  in the stratosphere and mesosphere in at least three ways. The first is strong convection in a thunderstorm cell that may carry tropospheric air into the lower stratosphere (Huntrieser et al. 2007). The second is by chemical effects of energetic electrons from the magnetosphere precipitated into the mesosphere and thermosphere from solar wind interactions with the magnetosphere or by whistler mode waves from lightning discharges (Rycroft 1973; Rodger et al. 2007). The third is local production of  $\text{NO}_x$  by sprites. The processes in sprites are in many ways similar to those in the aurora, although on smaller spatial and temporal scales, and can therefore be assumed to have at least local effects on middle-atmospheric  $\text{NO}_x$  chemistry. In the troposphere lightning is a significant source of  $\text{NO}_x$  (7 Tg/year), comparable to burning of biomass (8 Tg/year) and fossil fuels (22 Tg/year) (Schumann and Huntrieser 2007). Here a significant production channel of  $\text{NO}_x$  is through energetic neutral atom conversion where the heating of the atmosphere by the continuing current in the ion channel of the lightning stroke allows suprathreshold oxygen atoms to react directly with  $\text{N}_2$  (Balakrishnan and Dalgarno 2003). There is as yet no evidence for significant continuing currents in sprites or heating of the atmosphere beyond a few degrees, so this channel is not considered relevant for sprites in the mesosphere.

The effects of sprite discharges on the chemistry of the middle atmosphere have been addressed by several recent theoretical studies, based on independent streamer propagation models from a particle code, with inputs from spectral observations of sprites taken aboard a satellite (Enell et al. 2008), from a fluid code of streamer propagation (Sentman et al. 2008), or by a plasma code with a Boltzmann solver (Gordillio-Vazquez 2008). Here some results of the model by Enell et al. (2008) are summarized. As there is no unambiguous

evidence for a continuing current in sprites it is assumed that suprathermal atoms do not occur in sprites, although they may in the longer-lived blue jets and gigantic jets (Su et al. 2003). An additional point to be made here concerns experiments showing that sprite-like discharges (streamers) produce less  $\text{NO}_x$  per event than lightning-like (arc) discharges (Peterson et al. 2005). This is, however, only because arcs are more energetic than streamers. The chemical production rates of oxygen radicals, ozone, etc., per unit of dissipated energy can be much higher in streamers, which is the reason for using streamer technology for many energy efficient gas processing purposes (Winands 2007).

## 9.1 The Chemistry Code

The assumption that the ions and neutrals of the sprite plasma are not heated more than a few degrees above the background neutral temperature suggests the use of the Sodankylä coupled Ion-neutral Chemistry (SIC) model for estimating the  $\text{NO}_x$  production by sprites. SIC was originally developed for modelling of perturbations to the chemistry of the upper mesosphere and lower thermosphere induced by energetic charged particles precipitating from the magnetosphere, as inferred from observations by riometers, incoherent scatter radars, satellite-borne instruments and other measurements (Turunen et al. 1996). The present standard version of SIC6.x (Verronen 2005; Verronen (2006) is a time-dependent box model, where the coupled differential equations describing of the concentrations of the species are integrated for one altitude at a time. Almost 400 reactions of positive ions, negative ions and minor neutral species are taken into account. The reaction rate coefficients depend on neutral temperature and pressure but the reaction energetics are not treated explicitly. Ionization in SIC6.x is forced by sunlight (visible, UV and X rays), galactic cosmic rays, electrons and protons. Solar radiation and background density profiles are constant during one timestep of a complete altitude range update. Vertical transport, radiative transfer and background atmosphere altitude profiles are updated between the intervals. The profiles of the temperature and the neutral background species are taken from the MSISE-90 empirical model (Hedin 1991) and from tables given in Shimazaki 1984. Water vapour and  $\text{CO}_2$  are less well represented. For these we use fixed mixing ratios of 5 and 335 ppmv, respectively. The water mixing ratio affects  $\text{HO}_x$  production which is also important for ozone chemistry (Verronen et al. 2006). This issue merits further future studies. To estimate the chemical changes during a transient event, the SIC model is first run for a repeated full diurnal cycle in order to reach a quasi-steady state, which provides the reference background concentrations with which the event scenario is compared.

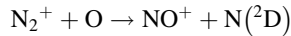
An updated version (SIC7.x) has been written for modelling sprites in the mesosphere (Enell et al. 2008). SIC7 includes 465 reactions of 29 negative ions, 45 positive ions and 50 neutral species. The most important excited states of  $\text{N}_2$ ,  $\text{O}_2$ ,  $\text{N}_2^+$ , N and O are now treated as separate species, and the reaction rate coefficients are functions of the electron temperature. An interface has been added to assimilate rates of excitation, ionization and dissociation of nitrogen and oxygen as obtained from other methods such as optical observations or discharge simulations.

In SIC7.x, production of nitric oxide in sprites occurs through three channels:

1. Dissociation of molecular oxygen and subsequent reactions of  $\text{O}(1\text{D})$  atoms such as  $\text{O}(1\text{D}) + \text{N}_2\text{O} \rightarrow 2\text{NO}$ .
2. Reactions of the long-lived excited state  $\text{N}_2(A^3\Sigma_u^+)$ . Its transition to the ground state, called the Vegard-Kaplan band, is highly forbidden and therefore the main loss is through chemical quenching (Vallance Jones 1974).



### 3. Ionization of $N_2$ :

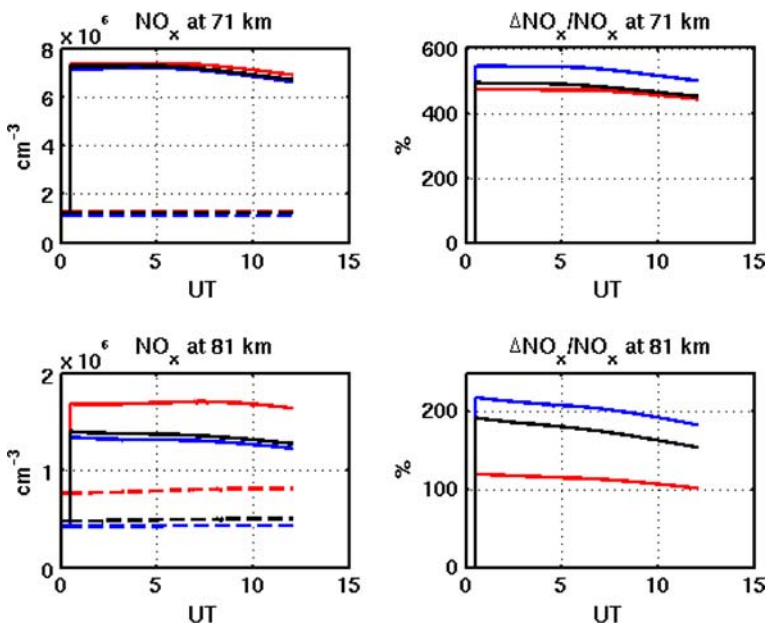


followed by reaction (2) and recombination of  $NO^+$  with electrons.

Reaction (3) is of less importance since  $N_2^+$  is quickly lost by charge exchange with  $O_2$ , and hence is not available to form either  $O^+$  or  $N(^2D)$  (Sentman et al. 2008).

### 9.2 Model Results on $NO_x$ Production in Sprites

Estimates of the enhancements of  $NO_x$  by sprites predicted by SIC7.x have been presented by Enell et al. (2008). The input rates of dissociation, excitation and ionisation are estimated from optical measurements of sprites made by a photometer on the FORMOSAT-2 satellite (Adachi et al. 2006). Since the measurements are not directly applicable because of lack of spatial and temporal resolution, they are instead used indirectly to estimate the electric field in the sprite streamers assuming that the photons primarily come from this region. These fields are then used to derive needed input rates for dissociation, excitation and ionization of nitrogen and oxygen using the Siglo cross sections ([www.siglo-kinema.com](http://www.siglo-kinema.com)). Figure 22 shows the production of NO (left) and corresponding relative enhancement (right) for a background atmosphere at 00:30 UT during a northern-hemispheric summer night (August 1–2). The estimated total production of  $NO_x$  in this scenario is around 5 times the background in the streamers. In comparison the model of Gordillio-Vazquez (2008) predicts NO and  $NO_2$  enhancements of a factor ten, which is in fairly good agreement given that these models present early efforts to study the phenomenon.



**Fig. 22** Left column:  $NO_x$  production by sprites. A sprite burst is imposed at a latitude of 60° N (black solid line), at 40° N (blue solid line) and at the equator (red solid line). Dashed lines show the unperturbed atmosphere. Right column: The relative increase above the control run concentrations (Enell et al. 2008)

Assuming a streamer reaching from 50–80 km altitude, the average NO production in a streamer is  $\sim 5 \times 10^{12} \text{ m}^{-3}$ . For a streamer of 10 m by 10 m cross section the volume of a streamer is  $\sim 3 \times 10^6 \text{ m}^{-3}$  and the total number of NO molecules produced in a streamer is  $\sim 1.5 \times 10^{19}$ , with an estimated average total production on the order of  $10^{23}$ – $10^{25}$  molecules per event, based on the dimensions of a typical sprite event (Gerken et al. 2000). By comparison, Sentman et al. (2008) estimated a production on the order of  $5 \times 10^{19}$  molecules per streamer, which is consistent with the above result.

Since  $\text{NO}_x$  is long-lived in the mesosphere with conversions between NO and  $\text{NO}_2$  taking place through photochemistry, the chemical significance to the atmosphere of sprite-produced  $\text{NO}_x$  is determined by transport out of the active region. The local effective buildup may be limited to a few hours, thus allowing large thunderstorm systems to accumulate up to 5 times the background concentration locally over a sprite-active cell of a storm. The global impact of sprites can be estimated from the global sprite occurrence rates of  $\sim 3/\text{min}$  (Ignaccolo et al. 2006) discussed in Sect. 6. This rate implies a global production of the order of  $10^{31}$  molecules/year, which is of the same magnitude as the minimum total  $\text{NO}_Y$  (i.e.  $\text{NO}_x$  plus stratospheric “reservoir species” such as  $\text{N}_2\text{O}$ ,  $\text{N}_2\text{O}_5$  and  $\text{HNO}_3$ ) produced by solar proton events during a quiet year, and two orders of magnitude below its production by cosmic rays (Vitt and Jackman 1996). The stratospheric production by oxidation of  $\text{N}_2\text{O}$  is on the order of  $10^{34}$  molecules/year (Vitt and Jackman 1996). From these considerations it follows that only over very intense thunderstorms can sprites lead to a significant impact upon the local density of  $\text{NO}_x$ . However, the global impact of sprites is diluted by their rarity; thus the preliminary studies presented here suggest that sprite-related processes are unlikely to play a significant role in the total  $\text{NO}_x$  budget of the stratosphere and mesosphere, although further studies are needed to consolidate our understanding of sprite chemistry.

### 9.3 Effects of Sprites on Ozone

Ozone is affected by  $\text{NO}_x$  through catalytic cycles. The maximum sensitivity of ozone to  $\text{NO}_x$  is in the stratosphere below 50 km at altitudes, which is a region only marginally affected by sprite streamers. On the other hand, dynamical time scales are longer at 50 km, thus allowing a longer accumulation time. Above about 60 km  $\text{HO}_x$  dominates the ozone catalytic cycles. The perturbations to ozone from  $\text{HO}_x$  reactions were found in the SIC model to be below the precision of the model. However, the impact of the sprite-produced NO on ozone was estimated to decrease the ozone concentration by a few percent in the 50–60 km altitude range for NO enhancements of 2–10 (Enell et al. 2008). Significant ozone perturbations from sprites are then unlikely, except for particularly sprite-active storms.

### 9.4 Comparisons with Observations

The chemical impact of sprites is further studied using the nighttime measurements of  $\text{NO}_2$  concentrations in the middle atmosphere given by Arnone et al. 2008a. The measurements are by the Michelson Interferometer for Passive Atmospheric Sounding (MIPAS) instrument (Fischer et al. 2007) on board the European Space Agency (ESA) ENVIRONMENTAL SATellite (ENVISAT). The measurements were retrieved with the new two-dimensional Geofit Multi-Target Retrieval algorithm (Niro et al. 2007) which enhances the detectability of small signals. The retrieved  $\text{NO}_2$  was correlated with intense active thunderstorms detected by the World Wide Lightning Location Network (WWLLN) (Jacobson et al. 2006).

The study found enhancements of  $\text{NO}_2$  above active thunderstorms, suggesting sprite-induced  $\text{NO}_x$  perturbations. The local enhancement increases with altitude from a few percent at 47 km to tens of percent at 60 km, thus being roughly consistent with SIC model estimates. Enhancements were not found on larger regional scales, as confirmed by studies based on data from the GOMOS instrument on ENVISAT (Rodger et al. 2008).

## 10 Perturbations to Atmospheric Dynamics

The possible occurrence of sprite-induced ozone changes raises the question of their role in the large scale atmospheric system. To what extent could sprites significantly affect climate and contribute to its variability? As seen in Sect. 9, the effect on the amounts of ozone from sprites seems to be low, even though they could have some effect for the more intense thunderstorms. The transport of  $\text{NO}_x$  produced by the sprites in the middle atmosphere is effective, and the chemicals are dispersed in both the meridional and the zonal directions within some hours to days, depending on the location of the storm. In this section we study the effect of different magnitudes of ozone perturbations to the middle atmosphere to assess a possible effect due to sprites in a climatological mean.

The extent to which even the small ozone perturbations estimated in the previous section could impact climate resides in the non-linear behaviour of the atmosphere. The circulation of the middle atmosphere is fundamentally determined by the temperature gradient between the tropics and the poles, and by the dispersion of planetary waves propagating upwards from the troposphere. The temperature gradient between the tropics and the poles and the Coriolis force give rise to zonal winds, which are easterly in the winter hemisphere and westerly in the summer hemisphere. In the meridional plane, the circulation consists of stratospheric air rising from the tropics and sinking at mid-latitudes, and a mesosphere summer hemisphere to winter pole flow accompanied by air rising in the summer pole and sinking into the troposphere in the winter pole (termed the residual circulation). This circulation is driven by upward propagating planetary waves in the winter hemisphere which break in the stratosphere, release their westerly momentum and thus reduce the easterly zonal winds (Christiansen 2001). Since the propagation of the planetary waves is affected by the local strength of the zonal winds, it may be affected by localized ozone perturbations in the stratosphere. The waves cannot propagate in a westerly zonal wind and they may therefore not penetrate through the zero wind line which extends vertically at the equator from the tropopause throughout the middle atmosphere. The position of this zero wind line has a proven effect on the pattern of the planetary wave propagation (e.g., Gray 2003). The dynamics also control the downward propagation of stratospheric temperature anomalies into the troposphere (Christiansen 2001), which imply a possible impact of stratospheric ozone changes onto the tropospheric climate.

Depending on the magnitude of the ozone perturbation produced by sprites, the propagation of planetary waves and the Brewer-Dobson circulation could be affected. Several model experiments were performed to assess the strength of the ozone perturbation needed to produce a significant change in the stratospheric dynamics and the extent to which this change propagates down to the troposphere.

### 10.1 The Methodology

The study was performed on two atmospheric models. The stratospheric impact of different magnitudes of ozone perturbations was assessed using the Stratosphere-Mesosphere Model

(SMM) from the UK Met Office (Austin and Butchart 1992; Berg et al. 2007). The model has a vertical resolution of 0.125 hPa, or about 2 km, ranging from 314 hPa to 0.03 hPa, i.e. from the upper troposphere to the lower mesosphere. The horizontal resolution is  $5^\circ \times 5^\circ$  on a global grid. To simulate the troposphere we introduced transient daily geopotential heights from the NCEP/NCAR reanalysis data at the lower boundary (Berg et al. 2007). This solution provides a realistic and naturally variable planetary wave forcing to the model.

The SMM simulations were then coupled to an atmospheric global circulation model (AGCM) to assess a possible tropospheric response. The AGCM used is cycle 14 of the ARPEGE model from MeteoFrance (Déqué et al. 1994; Christiansen et al. 1997). The ARPEGE model has 41 vertical levels with a high resolution of 1–2 km in the low stratosphere but only few levels in the high stratosphere to low mesosphere. The horizontal resolution is the same as for the SMM, i.e.  $5^\circ \times 5^\circ$ . The oceans are in a prescribed climatological state which prevents long delays in the response to the perturbations compared to having a coupled ocean in the model. Neither of the models has active chemistry. While the SMM was run with a full seasonal cycle, the AGCM was run in perpetual January mode to save computing time, therefore focusing on the Northern hemisphere winter season. In fact, in this season the coupling between the stratosphere and the troposphere is the largest, through the propagation of the planetary waves.

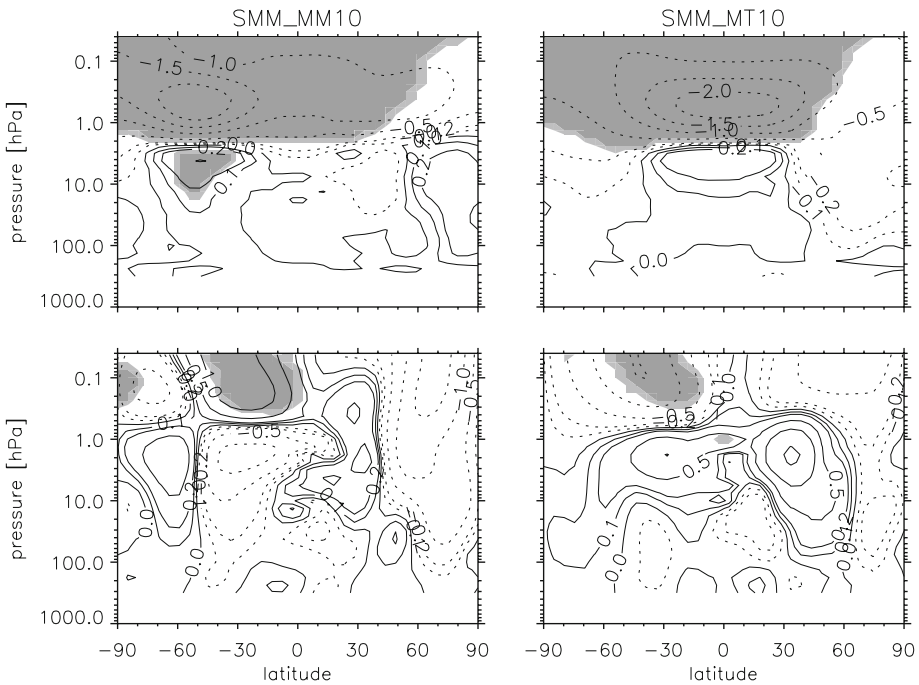
Since sprites have typical scales of tens of km, they are subgrid phenomena and cannot be modelled individually into standard global circulation models. As a first approximation, a climatological mean approach was thus adopted and sprites were assumed to be concentrated over the tropics and the midlatitudes where most thunderstorms occur. The experiments were then performed downscaling from very large magnitudes (50%) to lower ones, to find how large an ozone perturbation must be to produce a significant response in the model stratosphere and troposphere.

The SMM stratosphere was perturbed at the tropics ( $-30^\circ$  to  $+30^\circ$  latitude) and at the mid-latitudes ( $+/-30$  to  $70^\circ$  latitude) with zonally uniform ozone reductions of 5, 10, and 20% centred on the stratopause (10–0.1 hPa, i.e. about 30–80 km). Each experiment extended over 20 years with a full seasonal cycle, thus allowing statistical analyses to be performed. The perturbations centred at the stratopause were then repeated with the ARPEGE AGCM for ozone reductions of 10, 20 and 50%. The ARPEGE AGCM experiments were run in perpetual winter and extended over 600 days, which is roughly equivalent to 20 years assuming a decorrelation time of about 1 month.

In Sect. 9, the largest chemical impact of sprites was found to be associated with very intense sprite-producing thunderstorms which have timescales of several hours. An additional study of the sensitivity of the SMM stratosphere to repeated impulsive perturbations was thus carried out, introducing a time-dependency to the tropical stratopause perturbation experiment. In this case, the perturbation was introduced only for 1 day every 2 days, one every 4 days, every 8 days and every 16 days.

## 10.2 Results

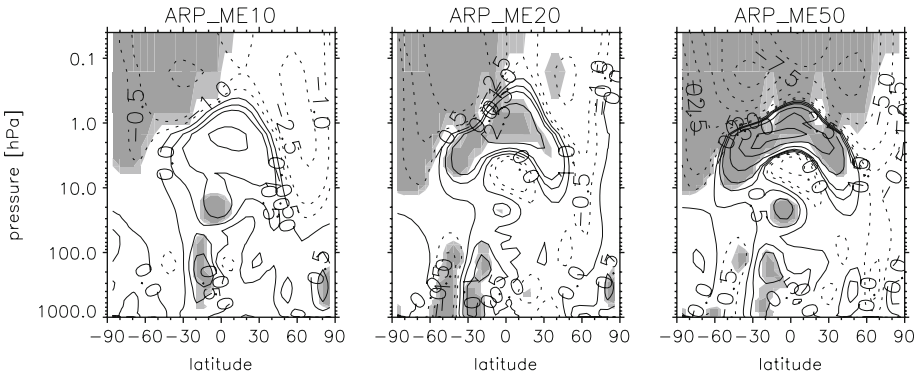
Figure 23 shows the results of the stratopause experiments (10–0.1 hPa, i.e. about 30–80 km) performed on the SMM at the tropics and at midlatitudes. In this figure, and in the following figures, the statistical significance levels of 90 and 95% are shaded. The testing was performed using the Student's t-test of means for mean values over each winter for the SMM, and for 30 day periods for the ARPEGE AGCM. Induced by the perturbations, both the mid-latitudinal and tropical regions show the largest negative response of the temperature



**Fig. 23** The SMM response in temperature (top) and zonal winds (bottom) to the stratopause 10% ozone reduction experiments at the midlatitudes (left) and tropics (right). The figure shows the difference between the perturbed and the control experiments. Responses at the 90% and 95% significance levels are shaded in grey

field within the perturbed region (top panels). Directly below the perturbed region there is a significant warming: this warming is due to absorption of the increased UV radiation that was not absorbed by the upper layers because of the ozone reduction. A 10% ozone decrease causes a temperature change which is significant for most of the latitudinal region at and above the stratopause, thus extending beyond the perturbed region through a dynamical response. The resulting change in the zonal wind (bottom panels) is weak and only significant in the lower mesosphere and around the stratopause. The changes follow the thermal wind approximation, i.e. that they are a direct result of the change in the temperature gradient. Inspection of the net heating rates (not shown) confirms a dynamical response (dynamical heating—see, e.g., Christiansen et al. 1997), with descending air confined in the perturbed regions. At both sides of these regions of dynamical heating are regions of dynamical cooling, i.e. the air is ascending. The statistical significance of the changes in the dynamical heating rates shows that not all of the changes are of a radiative nature. The patterns of the changes are consistent for all perturbation magnitudes, but the regions of significant change shrink with lower magnitudes. Downscaling from large (20%) to weak (5%) perturbations leads to a very confined local response.

The ARPEGE AGCM responds in a very similar way to the SMM in the middle atmosphere. Figure 24 shows the zonal wind response of the ARPEGE perturbation experiments performed at the tropical stratopause and with 10, 20 and 50% ozone reduction. Small regions of significant change appear in the tropospheric wind response which is roughly consistent for different perturbation magnitudes. Zonal winds show a significant shift toward the equator of the southern subtropical jet stream, which is



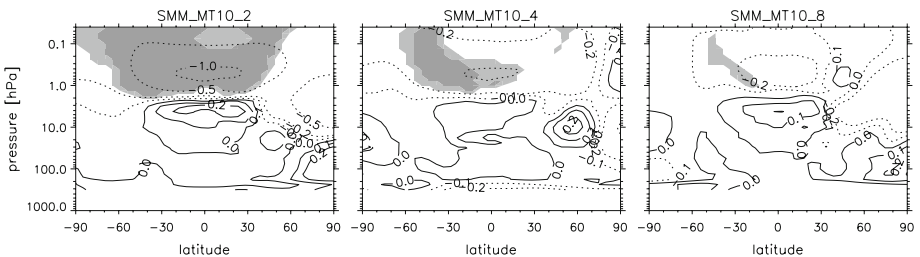
**Fig. 24** The ARPEGE response in zonal winds to the tropical stratopause ozone perturbation of magnitude 10%, 20% and 50% (left to right). The figure shows the difference between the perturbed and the control experiments. Responses at the 90% and 95% significance levels are shaded in grey

consistent also for the tropical case (not shown), even though the strength of the shift varies somewhat. Interestingly, the largest regions of statistical significant response are found in the 20% magnitude and are almost lost in the 50% case, showing a non-linearity in the propagation of the atmospheric response into the troposphere.

The clustering of sprite activity above intense thunderstorms was mimicked through the impulsive perturbation experiments. Figure 25 shows the temperature response for the tropical stratopause experiments with 10% ozone reductions applied for 1 day every 2 days, every 4 and every 8 days (top to bottom panels). The magnitude of the statistically significant response scales almost linearly with the average magnitude of the perturbation, disappearing below the experiment sensitivity with 1 day of 10% ozone reduction every 8 days. Below this limit, the response fades away, indicating that the magnitude of the perturbation is too small to induce a statistically significant response.

### 10.3 Discussion

A general result for both the middle atmosphere model and the GCM is that, in the stratosphere, the local changes in both temperature and zonal wind scale linearly with the magnitude of the perturbation in sunlit regions, for the perturbation magnitudes of a few to tens of percent decrease in ozone. The tropospheric response is not as predictable, with the 20% magnitude experiment leading to larger tropospheric responses than the 50% case.



**Fig. 25** The SMM temperature response to the 10% tropical stratopause ozone perturbation (as in Fig. 23 top-right) applied impulsively for 1 day every 2 days, one every 4 and one every 8 days (left to right). Notice that only the response at the 95% significance level is shown

This suggests a competing action of different radiative and dynamical processes, with the radiative drive taking over once the magnitude of the perturbation increases above a certain level. Dynamical processes are also at play in the winter hemisphere: especially near the Northern pole, the strongly disturbed temperature field makes it difficult for most perturbations to lead to a coherent statistically significant response. However, impulsive perturbations were found to be able more easily to alter the triggering of sudden stratospheric warming, leading to large impulsive temperature changes and possible propagation down to the troposphere.

Even accounting for these non-linear responses, the smallest perturbation needed by the model for any significant response (a zonally uniform stratopause 10% ozone reduction applied to the whole tropics for 1 day every 8 days) is far larger than the upper limit estimates deduced in the previous section (i.e. a sporadic 1% ozone reduction above very intense sprite-producing thunderstorms over a 300 km × 300 km region). These weak perturbations of the ozone concentrations are not likely to affect the tropospheric climate in any significant way through a radiative/dynamical pathway as shown by the model experiments. However, intense sprite activity is likely to cause a weak local temperature change and possibly lead to a transient perturbation of the local dynamics which cannot be resolved by the adopted global models.

## 11 Discussion

The past few years have seen a leap forward in our understanding of electric discharges in the mesosphere, both from activities of European research groups primarily reported on here and from the wider international community. However, while some areas of sprite physics are better understood, there are also significant questions raised that deserve attention in the years to come.

One question concerns the perturbation to the atmosphere by sprites. This issue is raised by the observation of infrasound waves from sprites, observed at up to 1000 km distance. The significance of the observation is its relation to the energy input to the atmosphere. Estimates in the past have been based on observations of the optical intensity of sprites in various spectral bands. They have varied by several orders of magnitude, with up to ~1 GJ for large events (Heavner 2000) and later the reduced estimates of 1–10 MJ (Sentman and São Sabbas 2002). The infrasound observations suggest large energy inputs, of the order of 0.4–40 GJ (Farges et al. 2005).

The question of energy input is related the level of ionization in sprites and the electric currents which they carry. Are sprites formed primarily a result of the excitation of atmospheric constituents by a modest density of electrons, perhaps released from negative ions, or is there significant ionization in sprites? Two regions of electric activity are involved in sprite generation; one is the region of sprites in the mesosphere and the other is the causative lightning discharge in the thunderstorm in the troposphere. Studies of the mesosphere are by necessity made by remote means and most commonly by observations of electromagnetic or optical signatures. The optical observations can distinguish between the two regions but electromagnetic detection techniques cannot. Thus there remains an ambiguity in electromagnetic studies concerning the source of observed signatures. Optical observations from the ground, on the other hand, are hampered by the absorption of emission lines associated with ionization that makes it difficult to estimate ionization rates in sprites. The observations obtained in the *EuroSprite* campaign of VLF transmitter signal perturbations suggest that sprites themselves create significant sources of enhanced

ionization at and below the lower boundary of the ionosphere. The conclusion is supported by studies of VLF radiation from lightning and sprites where efforts are made to distinguish between the two sources. Here the data suggest that sprites are sources of VLF radiation, which require electric currents to flow. This is further supported by the studies of others which combine optical data and electromagnetic measurements in the ELF range (Cummer et al. 2006b). In addition, imaging observations of interactions between branches of sprites clearly show the electric nature of these (Cummer et al. 2006a) and high-speed imaging show streamer tips, just as in laboratory discharges (McHarg et al. 2007).

We conclude, therefore, that sprites are electric discharges in the mesosphere with significant currents and significant perturbations to the atmosphere, as regards their detectability by conventional means. The quantitative estimation of all the perturbations, however, remains unsolved, in particular concerning heating of the neutral atmosphere, the plasma density of streamers and the electric currents which they carry.

A second aspect raised by the observations concerns the electric field established in the mesosphere following a lightning discharge. Although this field is perhaps the most important single entity to observe and understand in relation to sprite generation, our knowledge is rudimentary. It can only be estimated by a combination of modelling and observations of charge moment changes from ELF radio observations. In general, discussions on sprite generation consider the quasi-electrostatic field (ignoring the displacement current) from the discharge of a vertical dipole, simulating the cloud-to-ground discharge. However, observations of VLF radiation relating to sprite morphology clearly show the very complex nature of the problem, including a complex electric field source in the thunderstorms primarily due to horizontal intra-cloud and inter-cloud discharges. It is clear that we need to understand the complete electric fields produced by cloud-to-ground and intra/inter-cloud discharges. Very few attempts to model these latter fields have been made in the past (Valdivia et al. 1998; Cho and Rycroft 2001), but the advances made also in observations of the intra/inter-cloud component should allow further progress to be made in this area.

A third and final point relates to perturbations of the chemistry of the mesosphere. Here, models are available that can simulate a multitude of reaction paths and the challenge is to identify those which are the most important. This can only be done to the extent that the physics of the discharges is known. The simulations presented here have made assumptions about the sprite discharge, e.g., that there is no heating of the ambient atmosphere. While it is generally thought that little heating is associated with sprites, our infrasound observations indicate that this may not be true. At this point in our understanding we do not know how the neutral atmosphere is heated in sprite streamers or in the more diffuse region of the upper parts of carrot sprites. With this assumption, the studies of chemical perturbations to the concentrations of  $\text{NO}_x$  and ozone and the associated perturbations to atmospheric dynamics have shown that sprites have a definite local effect, perhaps regional effects, but insignificant global effects. The conclusions on the global effects are likely to remain unchanged in future, because of the limited rate of sprite events. However, more realistic models should be developed, and these may show unexpected effects.

Situated at 50–85 km altitude, the mesosphere is inaccessible to satellites, balloons and aircraft, and is difficult to probe by remote methods because of its tenuous character. As a result, detailed knowledge of the mesosphere is still lacking. Electric discharges in the mesosphere represent a spectacular, impulsive illumination of the region. With an in-depth understanding of the discharge process it will be possible to gain new insights into the composition and dynamics of this layer of the atmosphere, lying at the edge of space, at the boundary between the Earth's environment and the hostile space environment.



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## Appendix

### Appendix 1 Optical sprite observations over Europe during 2003 from Observatoire Midi Pyrénées

TLE#	TLE time	M	Flash time	Lat (North)	Long (East)	Peak I (A)	Distance (km)		
<i>Storm1</i>									
01.01	20030721	020514.630	1	02:05:14	7828741	45.705	3.075	−49.3	387.9
			1	02:05:14	8527730	45.916	2.700	23.1	389.0
01.02	20030721	020759.870	1	02:07:59	5380839	45.788	2.820	31.3	382.9
			2	02:07:59	7167443	45.796	2.755	32.3	380.6
01.03	20030721	020920.000	1	02:07:59	9124903	45.830	2.290	56.5	363.8
			1	02:09:20	153040	45.732	3.276	46.0	400.5
			2	02:09:20	641258	45.729	3.237	44.1	398.3
01.04	20030721	021151.750	1	02:09:20	1990054	45.887	2.940	32.6	397.2
			1	02:11:51	7182770	45.593	2.444	57.6	348.1
			1	02:17:00	8183461	45.739	3.135	45.8	393.9
01.05	20030721	021700.830	1	02:17:00	8808437	45.980	2.502	111.3	386.6
			1	02:19:46	2516176	46.837	3.680	185.8	516.2
01.06	20030721	021946.230	1	02:19:55	3601061	45.872	3.381	38.8	417.5
01.07	20030721	021955.430	1	02:22:59	4304597	46.097	3.758	28.5	455.6
01.08	20030721	022259.390	1	02:23:34	6286526	45.866	2.428	44.6	372.8
			1	02:23:34	6407980	45.959	2.421	16.5	381.4
01.09	20030721	022334.430	1	02:26:34	1805136	46.248	3.786	20.2	469.7
			2	02:26:34	1810964	46.239	3.694	22.4	464.3
			1	02:26:34	2855231	46.168	3.696	51.9	458.3
			1	02:26:34	4970698	46.266	3.318	39.7	448.5
01.10	20030721	022634.230	1	02:28:14	1960128	45.700	3.105	57.6	389.1
01.11	20030721	022814.270	1	02:28:14	1960128	45.700	3.105	57.6	389.1
01.12	20030721	022814.310	1	02:31:52	250349	45.554	2.841	36.7	363.4
01.13	20030721	023152.030	1	02:33:21	8277416	46.066	3.661	35.9	448.1
01.14	20030721	023321.870	1	02:33:47	3051369	45.730	3.314	24.6	402.3
01.15	20030721	023347.270	1	02:34:59	9947208	45.829	3.074	24.9	398.6
			1	02:35:00	1568612	45.902	2.588	49.3	382.9
01.16	20030721	023499.950	1	02:38:40	6642509	46.383	3.220	161.2	454.6
01.17	20030721	023840.590	1						

## Appendix 1 continued

TLE#	TLE time	M	Flash time	Lat (North)	Long (East)	Peak I (A)	Distance (km)		
		1	02:38:40	8186266	46.301	2.911	433.7		
		1	02:38:40	8477866	46.190	3.019	428.1		
01.18	20030721	024206.870	1	02:42:06	9145316	45.992	2.770	103.2	399.1
01.19	20030721	024332.470	1	02:43:32	3923382	46.475	4.387	42.8	519.4
01.20	20030721	024639.030	1	02:46:39	476445	45.482	3.364	46.6	385.0
			1	02:46:39	788618	45.394	3.570	33.7	390.1
01.21	20030721	025012.190	1	02:50:12	2396788	46.864	4.148	39.6	540.0
01.22	20030721	025038.790	1	02:50:38	7225762	45.604	3.342	25.9	393.5
01.23	20030721	025108.270	1	02:51:08	3103047	46.194	3.422	59.6	447.1
01.24	20030721	025703.630	1	02:57:03	4838543	45.741	3.456	17.5	410.7
01.25	20030721	025703.590	1	02:57:03	6687528	45.832	3.294	20.0	409.8
			1	02:57:03	9635451	46.242	3.215	18.0	441.6
01.26	20030721	030328.510	1	03:03:28	5468795	46.489	3.268	52.0	466.4
			1	03:03:28	6842424	46.166	3.102	40.1	429.6
01.27	20030721	030634.030	1	03:06:33	9252952	45.912	3.549	91.1	429.6
			1	03:06:34	850786	45.528	3.407	72.5	391.0
01.28	20030721	030925.550	1	03:09:25	1895247	45.500	3.862	17.2	415.0
Large delay			1	03:09:25	6027777	45.657	3.262	39.5	393.6
01.29	20030721	031334.230	1	03:13:34	581061	45.824	3.812	30.7	436.6
			1	03:13:34	2758955	45.824	3.393	50.9	414.1
<i>Storm2</i>									
02.01	20030722	215105.028	1	21:51:04	6203053	46.749	9.546	119.0	860.1
02.02	20030722	215430.294	1	21:54:30	3017605	45.064	5.604	102.7	503.8
02.03	20030722	215705.459	1	21:57:05	4447978	45.054	5.605	63.4	503.3
<i>Storm3</i>									
03.01	20030723	211132.326	1	21:11:32	2977297	44.598	3.866	60.4	357.9
03.02	20030723	213458.150							
03.03	20030723	215211.204	1	21:52:11	776952	44.990	5.039	33.6	460.2
03.04	20030723	215628.452	1	21:56:28	1965217	44.956	5.107	63.9	463.3
			2	21:56:28	2587845	44.976	5.053	27.3	460.5
03.05	20030723	215903.476	1	21:59:03	3271532	45.010	5.080	44.4	464.1
03.06	20030723	215903.496	1	21:59:03	3271532	45.010	5.080	44.4	464.1
03.07	20030723	220535.710	1	22:05:35	6789675	45.038	5.191	53.3	473.3
03.08	20030723	220906.796	1	22:09:06	6832650	45.047	5.127	25.4	469.2
03.09	20030723	220906.816	1	22:09:06	9047731	44.876	4.647	30.5	427.0
03.10	20030723	225102.258	1	22:51:02	992804	43.951	5.926	42.5	490.1
			1	22:51:02	2217239	44.106	5.728	75.5	478.1
03.11	20030723	231234.800	1	23:12:34	4376970	44.593	5.096	23.5	446.1
			1	23:12:34	8127213	44.367	5.452	54.8	464.4
			1	23:12:35	538850	44.223	5.762	46.7	483.9
03.12	20030723	232141.117	1	23:21:41	1091627	44.665	5.096	147.5	449.1
03.13	20030723	232141.257	1	23:21:41	1091627	44.665	5.096	147.5	449.1
			1	23:21:41	2725542	44.610	5.283	22.5	460.6

**Appendix 1** continued

TLE#	TLE time	M	Flash time	Lat (North)	Long (East)	Peak I (A)	Distance (km)
		1	23:21:41	4188476	44.463	5.592	23.5 478.3
03.14	20030724	002328.337	1 00:23:28	3197169	44.573	6.656	70.4 562.6
03.15	20030724	002827.485	1 00:28:27	4826391	44.686	6.711	59.5 570.4
03.16	20030724	003301.555	1 00:33:01	5385529	44.641	6.749	45.1 571.8
<i>Storm4</i>							
04.01	20030727	204619.521	1 20:46:19	3447626	46.514	9.783	17.2 865.8
			1 20:46:19	8040170	47.020	9.559	35.6 874.4
04.02	20030727	211355.814	1 21:13:55	6270131	47.488	9.913	49.2 922.6
			1 21:13:55	7977263	47.341	9.782	76.1 906.1
			1 21:13:55	9361352	47.073	9.527	42.7 875.0
<i>Storm5</i>							
05.01	20030820	201937.903	1 20:19:37	5350414	45.106	6.865	-14.5 596.9
<i>Storm6</i>							
06.01	20030822	212829.394	1 21:28:29	2436928	39.961	-0.761	25.3 344.2
06.02	20030822	212829.514	1 21:28:29	2436928	39.961	-0.761	25.3 344.2
06.03	20030822	213108.481	1 21:31:08	4962900	39.704	-1.095	55.4 378.2
			1 21:31:08	6075565	39.753	-1.268	68.9 376.6
06.04	20030822	213405.766	1 21:34:05	7828360	39.665	-0.787	56.6 377.0
06.05	20030822	213405.786	1 21:34:05	7828360	39.665	-0.787	56.6 377.0
06.06	20030822	213643.191	1 21:36:43	1933773	39.734	-1.073	113.8 374.4
06.07	20030822	214252.243	1 21:42:52	2564001	39.740	-0.735	99.6 368.0
<i>Storm7</i>							
07.01	20030824	211741.901	1 21:17:41	6923074	43.503	2.399	41.8 202.4
07.02	20030824	213742.359	1 21:37:42	2763870	43.443	2.543	129.9 212.1
			1 21:37:42	3696938	43.446	2.564	54.1 213.8
<i>Storm8</i>							
08.01	20030825	195647.448	1 19:56:46	8297394	42.538	11.219	26.5 917.5
			1 19:56:47	4590832	42.790	10.863	62.6 885.6
08.02	20030825	203037.810	1 20:30:37	8299507	42.477	11.120	144.8 910.3
08.03	20030825	203235.014	1 20:32:35	265978	42.381	11.266	71.5 923.6
08.04	20030825	205612.599	1 20:56:12	6021825	43.078	10.067	85.1 818.6
08.05	20030825	205955.386	1 20:59:55	3887244	42.775	9.884	88.9 806.1
			1 20:59:55	5520819	43.033	10.002	42.8 813.6
08.06	20030825	210830.763	1 21:08:30	7575068	43.064	9.407	62.5 765.1
08.07	20030825	211456.337	1 21:14:56	3165277	42.922	9.963	150.6 811.2
08.08	20030825	213517.357	1 21:35:16	9525907	42.445	9.682	25.2 793.7
			1 21:35:17	3605032	42.820	9.762	134.8 795.7
08.09	20030825	213517.397	1 21:35:17	3605032	42.820	9.762	134.8 795.7
08.10	20030825	213517.417	1 21:35:17	3605032	42.820	9.762	134.8 795.7
			1 21:35:17	4947733	43.233	9.763	47.1 793.3
08.11	20030825	213721.481	1 21:37:21	4824718	42.847	9.336	76.3 760.8
			1 21:37:21	5813983	43.028	9.659	34.2 785.8

**Appendix 1** continued

TLE#	TLE time		M	Flash time	Lat (North)	Long (East)	Peak I (A)	Distance (km)
<i>Storm9</i>								
09.01	20030828	201800.815	1	20:18:00	8216922	46.316	4.309	84.4 502.5
			1	20:18:00	9656683	46.389	3.955	45.5 490.0
09.02	20030828	225706.221	1	22:57:06	1293877	45.798	5.795	69.2 556.0
			1	22:57:06	3947627	46.047	6.091	55.1 590.3
			1	22:57:06	4725884	45.738	6.455	31.8 597.0
09.03	20030828	230324.974	1	23:03:24	9936912	45.348	5.509	110.6 511.5
09.04	20030828	231110.569						
09.05	20030828	231110.689						
09.06	20030828	231110.829	1	23:11:10	8468338	45.745	5.628	57.9 541.9
			1	23:11:11	2703137	45.984	6.063	24.5 584.7
			1	23:11:11	4065650	46.140	5.884	15.1 582.8
09.07	20030828	231355.134	1	23:13:55	1571086	45.450	5.592	124.2 522.6
09.08	20030828	231355.194	2	23:13:55	1938103	45.444	5.584	22.4 521.8
			1	23:13:55	5448105	45.744	5.698	18.1 546.4
09.09	20030828	231441.175	1	23:14:41	1379871	47.295	7.837	45.7 778.6
09.10	20030828	231655.799	1	23:16:55	5048005	45.492	5.652	117.6 529.0
09.11	20030828	231655.839	1	23:16:55	8033441	45.982	5.965	106.8 578.1
09.12	20030828	231655.859	1	23:16:55	8158742	45.466	6.049	41.0 554.9
09.13	20030828	231655.899	1	23:16:55	9034059	45.931	6.254	72.2 594.1
			1	23:16:55	9261566	46.060	6.052	36.1 588.5
09.14	20030828	231842.362	1	23:18:42	2239738	45.471	5.646	18.2 527.5
			2	23:18:42	3518902	45.487	5.704	68.0 532.3
09.15	20030828	231946.584	1	23:19:46	5570899	45.548	5.726	113.6 537.1
09.16	20030828	231946.984	1	23:19:46	9867420	45.710	6.345	112.9 588.0
09.17	20030828	231947.184	1	23:19:47	1936325	45.441	6.536	72.7 587.7
09.18	20030828	232154.569	1	23:21:54	5142768	45.579	5.765	83.7 541.5
09.19	20030828	232154.589	1	23:21:54	5142768	45.579	5.765	83.7 541.5
09.20	20030828	232154.789	1	23:21:54	7771318	45.832	5.866	20.0 562.8
09.21	20030828	232442.474	1	23:24:42	4254060	45.582	5.772	72.3 542.1
09.22	20030828	232920.382	1	23:29:20	2764177	45.298	5.882	71.1 534.7
09.23	20030828	232920.422	1	23:29:20	2764177	45.298	5.882	71.1 534.7
09.24	20030828	232920.662	1	23:29:20	6606870	45.735	6.295	60.0 585.9
09.25	20030828	232920.702	1	23:29:20	6606870	45.735	6.295	60.0 585.9
09.26	20030828	232920.782	1	23:29:20	6606870	45.735	6.295	60.0 585.9
09.27	20030828	233140.708						
09.28	20030828	233424.592	1	23:34:24	4886795	45.412	5.733	32.2 530.2
			1	23:34:24	7554713	45.684	5.887	30.9 555.6
09.29	20030829	000442.550	1	00:04:42	5532696	45.786	6.368	74.5 593.7
			1	00:04:42	7874671	45.657	6.650	56.0 606.3
09.30	20030829	004024.043	1	00:40:24	275916	45.530	5.821	46.0 542.6
09.31	20030829	004454.285	1	00:44:54	2034265	45.562	5.889	56.2 549.0
09.32	20030829	014111.170	1	01:41:11	973522	46.573	9.460	111.7 845.8

**Appendix 1** continued

TLE#	TLE time	M	Flash time	Lat (North)	Long (East)	Peak I (A)	Distance (km)
09.33	20030829 021601.574	1	02:16:01	5827200	45.964	6.051	129.7 582.7
09.34	20030829 021643.235						
09.35	20030829 023414.607	1	02:34:14	6231497	45.844	5.836	106.3 561.4
09.36	20030829 023414.747	1	02:34:14	7637664	46.006	6.027	112.0 583.6
		1	02:34:14	8716812	46.080	6.405	93.9 612.7
09.37	20030829 023519.989	1	02:35:19	9960544	46.643	9.691	45.6 865.2
		1	02:35:20	970295	46.783	9.897	25.9 885.9
08.12	20030825 214129.390	1	21:41:29	4038027	42.802	9.061	42.5 738.8
08.13	20030825 214129.450	1	21:41:29	5111275	42.968	9.497	52.6 773.0
08.14	20030825 214129.490	1	21:41:29	6144416	42.823	9.812	69.4 799.7
		1	21:41:29	6227234	43.301	9.555	39.1 776.3
08.15	20030825 214454.916	1	21:44:55	1214298	43.292	9.614	31.8 781.1
08.16	20030825 214454.936	1	21:44:55	1396312	43.058	9.324	32.8 758.4
08.17	20030825 215022.015	2	21:50:21	9421691	43.104	8.818	59.5 717.1
		1	21:50:22	252537	43.132	9.032	63.2 734.3
		1	21:50:22	1589191	42.941	9.666	34.8 786.9
08.18	20030825 215659.620	1	21:56:59	6192954	43.123	8.723	53.9 709.3
08.19	20030825 215659.720	1	21:56:59	7213260	43.152	9.256	47.7 752.4
08.20	20030825 220634.519	1	22:06:34	4835368	43.121	8.742	45.7 710.8
09.38	20030829 023826.394						
09.39	20030829 023922.816						
09.40	20030829 024110.641	1	02:41:10	6582007	46.061	6.073	76.2 590.0
09.41	20030829 025332.523	1	02:53:32	4942708	46.560	9.748	41.1 865.3
09.42	20030829 025636.888	1	02:56:36	8960817	46.299	6.291	68.2 618.6
09.43	20030829 025753.230	1	02:57:53	2270698	46.827	10.342	98.4 918.8
09.44	20030829 030525.885	1	03:05:25	8787280	46.508	9.954	41.8 877.5
09.45	20030829 031340.460	1	03:13:40	3181245	46.072	6.096	38.1 592.1
09.46	20030829 031946.970	1	03:19:46	9768514	46.823	10.229	19.6 910.8
		1	03:19:47	1633218	46.992	10.407	30.7 930.9
09.47	20030829 033552.860						
09.48	20030829 033552.900						
09.49	20030829 034419.996	1	03:44:20	214724	45.184	5.101	-141.5 474.7
09.50	20030829 034543.638	1	03:45:43	6640816	46.402	6.503	268.7 638.5
09.51	20030829 034543.678	1	03:45:43	6640816	46.402	6.503	268.7 638.5
09.52	20030829 034707.201						
09.53	20030829 034707.221						

The times (UT) of sprite events are shown with the corresponding geographic location (in degrees and thousandths of a degree), peak current I and multiplicity of the associated +CG lightning

Also shown is the distance from the cameras to the events. Transient Luminous Event (TLE) time is given in the following format: YYYYMMDD (column 2) HHMMSS.SSS (column 3). Flash time is in the format: HH:MM:SS (column 6) and tenths of microseconds (column 7). M is the multiplicity of the lightning flash

**Appendix 2** Optical sprite observations over Europe during 2005 from Observatoire Midi Pyrénées and Puy de Dôme (Armone et al. 2008b)

Event	Station	Time (UTC)	GPS accuracy	+CG location (lat, long)	+CG (kA)	Time +CG (UTC)	CMC (Ckm)
1	Pic	2005.07.19	OK 10 $\mu$ s	45°3626; 9°9975	70.7	00:46:12.4340157	1042
2	Pic	2005.07.29	~<1 s	45°0439; 2°2686	197.4	01:28:59.0650213	1387
3a	Pic	2005.07.29	~<1 s	45°0349; 2°2981	38.7	01:28:59.0771788	
3b	Pic	2005.07.29	~<1 s	44°9256; 2°4655	24.6	01:32:51.2459586	1460
4	Pic	2005.07.29	~<1 s	44°9824; 2°2419	35.5	01:32:51.3647629	
5	Pic	2005.07.29	~<1 s	45°2527; 2°6821	66.7	01:36:09.3574440	
6a	Pic	2005.07.29	~<1 s	45°2544; 2°6796	29.4	01:36:09.3938580	
6b	Pic	2005.07.29	~<1 s	44°9953; 2°6827	61.4	01:40:50.1954235	1020
7	Pic	2005.07.29	~<1 s	45°0610; 2°1906	115.6	01:40:50.3783621	
8a	Pic	2005.07.29	~<1 s	45°2205; 2°5627	8.1	01:43:52.5248687	
8b	Pic	2005.07.29	~<1 s	45°3520; 2°0880	25.5	01:43:52.3707276	
9	Pic	2005.07.29	~<1 s	45°2653; 2°7139	63.1	01:43:52.7369657	
10a	Pic	2005.07.29	~<1 s	45°3333; 2°7731	129.9	01:46:04.7391040	446
10b	Pic	2005.07.29	~<1 s	45°2719; 2°6210	75.2	01:48:08.2671118	2239
11a	Pic	2005.07.29	~<1 s	45°3221; 2°6180	16.2	01:48:08.4187157	
			~<1 s	45°2976; 2°0029	27.8	01:48:08.5516591	761
			~<1 s	46°2103; 3°0672	49.8	01:54:35.1570309	
			~<1 s	45°0637; 2°8356	71.9	01:54:35.3039434	
			~<1 s	45°4945; 2°6584	31.0	01:54:35.3051241	1106
			~<1 s	45°3311; 2°7704	99.7	01:54:38.0782815	
			~<1 s			01:54:38.4447377	
			~<1 s			02:04:10.2559617	

## Appendix 2 continued

Event	Station	Time (UTC)	GPS accuracy	+CG location (lat, long)	+CG (kA)	Time +CG (UTC)	CMC (Ckm)
11b	Pic	2005.07.29	~<1 s	45°5442; 2°5895 45°4436; 2°5321 45°2514; 2°5886	66.2 or 32.9 or 53.6	02:04:10.3429030 02:04:10.3434299 02:04:10.3684963	
12	Puy	2005.08.11	~<1 s	43°9625; 6°8845	43.0	22:43:26.5979664	945
13	Puy	2005.08.11	~<1 s	43°5969; 6°7493 43°8537; 7°0086	80.6 31.6	23:01:31.9760132 23:01:32.0703604	1171
14	Puy	2005.08.11	~<1 s	43°7520; 6°7169	50.7	23:11:50.1970008	680
15	Pic	2005.08.16	OK 10 $\mu$ s	41°9948; 0°8384	24.5	23:06:42.1824716	
16	Pic	2005.09.07	OK 10 $\mu$ s	43°8377; 6°2490	57.6	02:14:08.3374678	1483
17a	Puy	2005.09.09	~<1 s	43°8410; 5°4698 43°8933; 5°4027	47.4 25.6	01:08:08.9138779 01:08:08.9924877	
17b	Puy	2005.09.09	~<1 s	43°7617; 6°1867	114.4	01:08:09.0845451	563
18	Puy	2005.09.09	~<1 s	43°6843; 5°8300	177.9	03:02:12.6417582	1282
19	Puy	2005.09.09	~<1 s	43°8566; 7°1371 43°9189; 7°0820	170.3 81.2	03:45:42.3507999 03:45:42.3886371	
20	Pic	2005.09.09	OK	44°6450; -1°4153	37.2	20:25:42.2929545	582
21	Puy	2005.09.09	~<1 s	44°8078; -0°8396 44°8404; -1°0773	98.0 43.7	20:25:42.2929545 20:52:24.5429706	520
22	Pic	2005.09.09	OK 10 $\mu$ s	45°0097; -1°0484	74.5	21:18:39.2203371	
23	Pic	2005.09.26	OK 10 $\mu$ s	40°3565; 3°2920	77.4	01:06:35.6185602	1228
24	Pic	2005.09.26	OK 10 $\mu$ s	40°5905; 3°2933	89.5	01:09:43.5500720	3673
25	Pic	2005.09.26	OK 10 $\mu$ s	40°4105; 3°8076 40°5277; 3°6858	41.2 3.6	01:14:39.4762022 01:14:39.7407460	507
26a	Pic	2005.09.26	OK 10 $\mu$ s	40°5682; 3°9882	70.7	01:22:03.8229370	
26b	Pic	2005.09.26	OK 10 $\mu$ s	40°6368; 3°3994	19.2	01:22:03.9696297	973

## Appendix 2 continued

Event	Station	Time (UTC)	GPS accuracy	+CG location (lat, long)	+CG (kA)	Time +CG (UTC)	CMC (Ckm)
27	Pic	2005.09.26	OK 10 $\mu$ s	40°4839; 4°0563 40°5396; 3°4241 40°5677; 3°0873	28.4 56.1 76.5	01:25:38.8947767 01:25:38.9968106 01:25:38.9968856	
28a	Pic	2005.09.26	OK 10 $\mu$ s	40°5257; 4°0132	138.4	01:33:43.2475989	
28b	Pic	2005.09.26	OK 10 $\mu$ s	40°8745; 3°4924 40°5765; 3°3765	60.6 84.0	01:33:43.3480423 01:33:43.3595888	2578
29a	Pic	2005.09.26	OK 10 $\mu$ s	40°4881; 3°9481	12.9	01:39:09.6137574	358
29b	Pic	2005.09.26	OK 10 $\mu$ s	40°6594; 3°6072 40°7859; 3°6114	26.4 8.2	01:39:09.8234962 01:39:09.8480144	
30	Pic	2005.09.26	OK 10 $\mu$ s	40°5616; 3°6740	112.9	01:49:15.2933991	1677
31	Pic	2005.09.26	OK 10 $\mu$ s	40°5265; 3°9523	126.8	01:54:17.4430634	2389
32	Pic	2005.09.26	OK 10 $\mu$ s	No +CG found	No +CG found	No +CG found	1704
33	Pic	2005.09.26	OK 10 $\mu$ s	No +CG found	No +CG found	No +CG found	
34	Pic	2005.09.26	OK 10 $\mu$ s	40°4284; 3°9302	10.4	02:49:52.228621	801
35	Pic	2005.09.26	OK 10 $\mu$ s	40°2957; 4°0049	10.7	02:53:57.9171687	
36a	Pic	2005.09.26	OK 10 $\mu$ s	40°3054; 3°7981	60.8	03:00:42.5285112	1413
36b	Pic	2005.09.26	OK 10 $\mu$ s	40°5265; 3°5147 41°0074; 3°9667	97.1 5.0	03:00:42.7187269 03:00:42.7450610	
37a	Pic	2005.09.26	OK 10 $\mu$ s	40°4541; 4°1782	26.4	03:06:19.3339677	506
37b	Pic	2005.09.26	OK 10 $\mu$ s	40°8201; 3°3885	85.8	03:06:19.7975058	936
38	Pic	2005.09.26	OK 10 $\mu$ s	40°3902; 3°9138	105.5	03:10:29.1952754	1346
39a	Pic	2005.09.26	OK 10 $\mu$ s	40°6580; 3°9072	25.0	03:12:41.5266687	
39b	Pic	2005.09.26	OK 10 $\mu$ s	40°4573; 3°6664	27.7	03:12:41.6383932	1135
40	Pic	2005.09.26	OK 10 $\mu$ s	40°8010; 3°8467	274.5	03:23:03.9238399	
41	Pic	2005.09.26	OK 10 $\mu$ s	40°5872; 4°0318	64.3	03:27:22.6788673	1101



## Appendix 2 continued

Event	Station	Time (UTC)	GPS accuracy	+CG location (lat, long)	+CG (kA)	Time +CG (UTC)	CMC (Ckm)
42	Pic	2005.09.26	OK 10 $\mu$ s	40°1432;5°8353	4.5	03:38:57.2366405	4591
43	Puy	2005.10.06	~<1 s	No +CG found	No +CG found	No +CG found	
44	Puy	2005.10.06	~<1 s	No +CG found	No +CG found	No +CG found	
45	Puy	2005.10.06	~<1 s	No +CG found	No +CG found	No +CG found	
46	Pic	2005.10.23	OK 10 $\mu$ s	40°6406;3°7546	41.5	18:59:23.9424131	1233
47	Pic	2005.11.11	~<1 s	40°0550;3°1835	26.8	00:56:46.5545095	
				39°9641;2°9490	97.0	00:56:46.5553589	
48	Pic	2005.11.11	~<1 s	40°1907;2°8985	28.3	01:11:31.1396087	522
49	Pic	2005.11.11	~<1 s	40°6097;4°6513	7.1	02:34:13.4989533	
						02:34:13.4989533	
50a	Pic	2005.11.11	~<1 s	40°4398;3°8508	10.1	02:43:46.6126729	
50b	Pic	2005.11.11	~<1 s	40°5305;3°4553	30.4	02:43:46.7760099	1281
				40°3586;3°4009	45.3	02:43:46.8189332	
50c	Pic	2005.11.11	~<1 s	40°3830;2°9583	89.0	02:43:47.0418297	1258
51a	Pic	2005.11.11	~<1 s	40°7560;3°6362	41.0	02:56:13.8272081	
51b	Pic	2005.11.11	~<1 s	40°5668;3°3505	17.5	02:56:13.8463774	
51c	Pic	2005.11.11	~<1 s	40°6070;2°9842	30.1	02:56:13.9605701	
52	Pic	2005.11.12	~<1 s	37°803; -0°1312	13.2	02:21:38.3739785	3421
53	Pic	2005.11.12	~<1 s	No +CG found	No +CG found	No +CG found	
54	Pic	2005.11.12	~<1 s	No +CG found	No +CG found	No +CG found	
55	Pic	2005.11.12	~<1 s	38°5832;0°4316	27.1	04:27:53.9063009	2043
56	Pic	2005.11.12	~<1 s	38°5189;0°6066	108.0	04:31:19.3083936	1448
57	Pic	2005.11.12	~<1 s	No +CG found	No +CG found	No +CG found	
58	Pic	2005.11.12	~<1 s	38°5522;1°5099	13.8 or	05:02:02.8564192	2821
				38°3828;1°2003	72.4	05:02:02.9173652	

## Appendix 2 continued

Event	Station	Time (UTC)	GPS accuracy	+CG location (lat, long)	+CG (kA)	Time +CG (UTC)	CMC (Ckm)
59	Pic	2005.11.15	OK 10 $\mu$ s	No +CG found	No +CG found	No +CG found	2484
60	Pic	2005.11.15	OK 10 $\mu$ s	No +CG found	No +CG found	No +CG found	
61	Pic	2005.11.15	OK 10 $\mu$ s	40°6283;2°0396	2.8	22:12:43,4459992	1635
62	Pic	2005.11.15	OK 10 $\mu$ s	No +CG found	No +CG found	No +CG found	
63	Pic	2005.11.15	OK 10 $\mu$ s	No +CG found	No +CG found	No +CG found	
64	Pic	2005.11.15	OK 10 $\mu$ s	No +CG found	No +CG found	No +CG found	

The time of sprite events are shown with the corresponding geographic location, peak current and time of the associated +CG lightning. Also shown is the charge moment change (CMC) calculated from Nagyenk Schumann resonance data. For some events more than one +CG could be associated with the sprite. In a few cases no +CG could be identified

**Appendix 3** Optical sprite observations over Europe during 2006 from Observatoire Midi Pyrénées

Event	Time	(UTC)	Accuracy GPS	Time +CG (UTC)	+CG location (lat, long)	I +CG (kA)
1	2006.07.03	022723.298	~20 ms	No +CG	No +CG found	
2	2006.07.03	023456.918	~20 ms	023456.877	44.308700; -0.574400	25.6
3	2006.07.03	023723.748	~20 ms	023723.754	44.976900; -1.360700	40.7
				023727.831	45.605400; -2.529100	8.9
4	2006.07.03	024255.164	~20 ms	024255.068	45.069900; -0.818000	31.0
				024255.288	44.910000; -1.184400	32.9
5	2006.07.03	024608.421	~20 ms	024608.389	45.629300; -2.016500	7.8
6	2006.07.03	025238.185	~20 ms	025238.176	45.299500; -1.073700	22.6
7	2006.07.04	204606.591	~20 ms	204606.598	44.350600; -0.808200	81.8
8	2006.07.04	223651.838	OK 10 $\mu$ s	223651.733	47.247000; -0.139000	55.6
9	2006.07.04	225153.730	OK 10 $\mu$ s	225153.733	46.937100; -0.848000	52.8
10	2006.07.05	000436.446	OK 10 $\mu$ s	000436.299	45.228400; -1.917100	19.1
				000436.446	45.269400; -1.966500	38.2
11	2006.07.05	001642.461	OK 10 $\mu$ s	001642.422	45.578900; 2.084900	76.5
12	2006.07.05	003843.416	OK 10 $\mu$ s	003843.522	45.572500; 2.103000	39.3
				003843.524	45.459600; 2.139000	48.8
				003843.596	45.243000; 2.242500	29.6
13	2006.07.05	030531.032	~500 $\mu$ s	030531.017	45.138300; 2.603800	85.3
14	2006.07.05	224233.711	OK 10 $\mu$ s	No +CG	No +CG found	
15	2006.07.21	220615.805	OK 10 $\mu$ s	220615.521	43.473600; 0.743600	33.5
16	2006.07.22	014657.000	~10 ms	014657.004	44.731200; 0.999900	45.2
17	2006.07.22	020132.056	~10 ms	020132.056	44.871800; 1.169100	61.8
18	2006.07.27	011317.209	OK 40 $\mu$ s	011317.168	45.957100; 1.076500	51.0
19	2006.07.27	013657.253	~15 ms	013657.202	46.145900; 1.519200	138.4
20	2006.07.27	022418.582	~4 ms	022418.533	47.612600; 2.061200	25.5
				022418.570	47.622400; 1.970000	99.3
				022418.880	47.954800; 1.745600	56.1
21	2006.07.27	023317.129	OK 10 $\mu$ s	023317.109	47.381900; 1.038200	82.0
				023317.235	47.618900; 1.182800	87.9
				023317.454	47.716200; 1.177600	33.4
				023317.604	47.713200; 1.730000	36.8
				023317.928	47.868200; 1.746600	36.8
				023318.096	48.070100; 1.949700	56.9
				023318.167	47.557800; 1.804800	32.3
				023318.551	47.867900; 1.594000	63.4
22	2006.07.27	023650.569	OK 10 $\mu$ s	023650.522	47.461200; 1.125200	136.0
				023650.633	47.634300; 1.209400	104.2
				023650.740	47.542700; 1.207200	29.5
				023650.807	47.698500; 1.726700	31.6
23	2006.08.17	222944.167	~20 ms	222944.159	44.778700; 4.351300	55.2
				222944.729	49.522400; 8.051100	7.2
				222944.855	48.673700; 6.562400	25.8

**Appendix 3** continued

Event	Time	(UTC)	Accuracy GPS	Time +CG (UTC)	+CG location (lat, long)	I +CG (kA)
24	2006.08.17	223400.272	OK 20 $\mu$ s	223400.221	44.907700; 4.339200	57.2
25	2006.08.17	225058.941	$\sim$ 40 ms	225058.535	48.368800; 7.249700	24.0
				225058.918	45.196300; 4.842500	100.0
26	2006.08.17	225333.480	$\sim$ 40 ms	225333.431	45.134000; 4.766600	44.2
				225333.540	45.070400; 4.482800	19.1
27	2006.08.17	225846.222	$\sim$ 1 ms	225846.219	45.136200; 4.421300	115.3
				225846.648	44.968800; 4.819300	27.0
28	2006.08.17	230049.125	$\sim$ 1 ms	230049.141	45.377700; 4.888600	141.4
29	2006.08.17	233807.755	$\sim$ 1 ms	233807.719	44.716400; 5.177600	29.5
				233808.102	45.325800; 5.255600	68.9
30	2006.08.17	234732.476	$\sim$ 6 ms	234732.373	44.741200; 5.268800	45.0
				234732.603	45.164800; 5.211800	73.8
				234732.677	45.166300; 5.044000	12.3
				234732.728	45.587000; 4.846200	34.8
				234732.933	45.990600; 4.972000	25.1
31	2006.08.17	235316.611	OK 20 $\mu$ s	235316.610	44.842200; 5.379800	62.2
				235316.746	45.387000; 5.085800	35.1
				235316.802	45.495000; 5.129200	16.1
				235316.922	45.856600; 4.977100	47.3
32	2006.08.17	235715.386	OK 50 $\mu$ s	235715.262	44.692000; 5.449400	40.7
				235715.489	45.283000; 5.325400	139.6
				235715.545	45.468800; 5.189200	107.4
33	2006.08.18	001427.977	OK 1 ms	001427.837	44.735300; 5.695900	31.9
				001427.522	44.516800; 5.605700	27.8
				001427.912	44.961800; 5.841300	40.4
				001428.020	45.387300; 5.545800	118.9

The time of sprite events are shown with the corresponding geographic location, peak current and time of the associated +CG lightning. For some events more than one +CG could be associated with the sprite. In a few cases no +CG could be identified

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